

# MACHINE LEARNING RECONSTRUCTION OF EPOCH OF REIONIZATION BUBBLES



Institut Spatial de McGill



McGill Space Institute

**CIFAR**

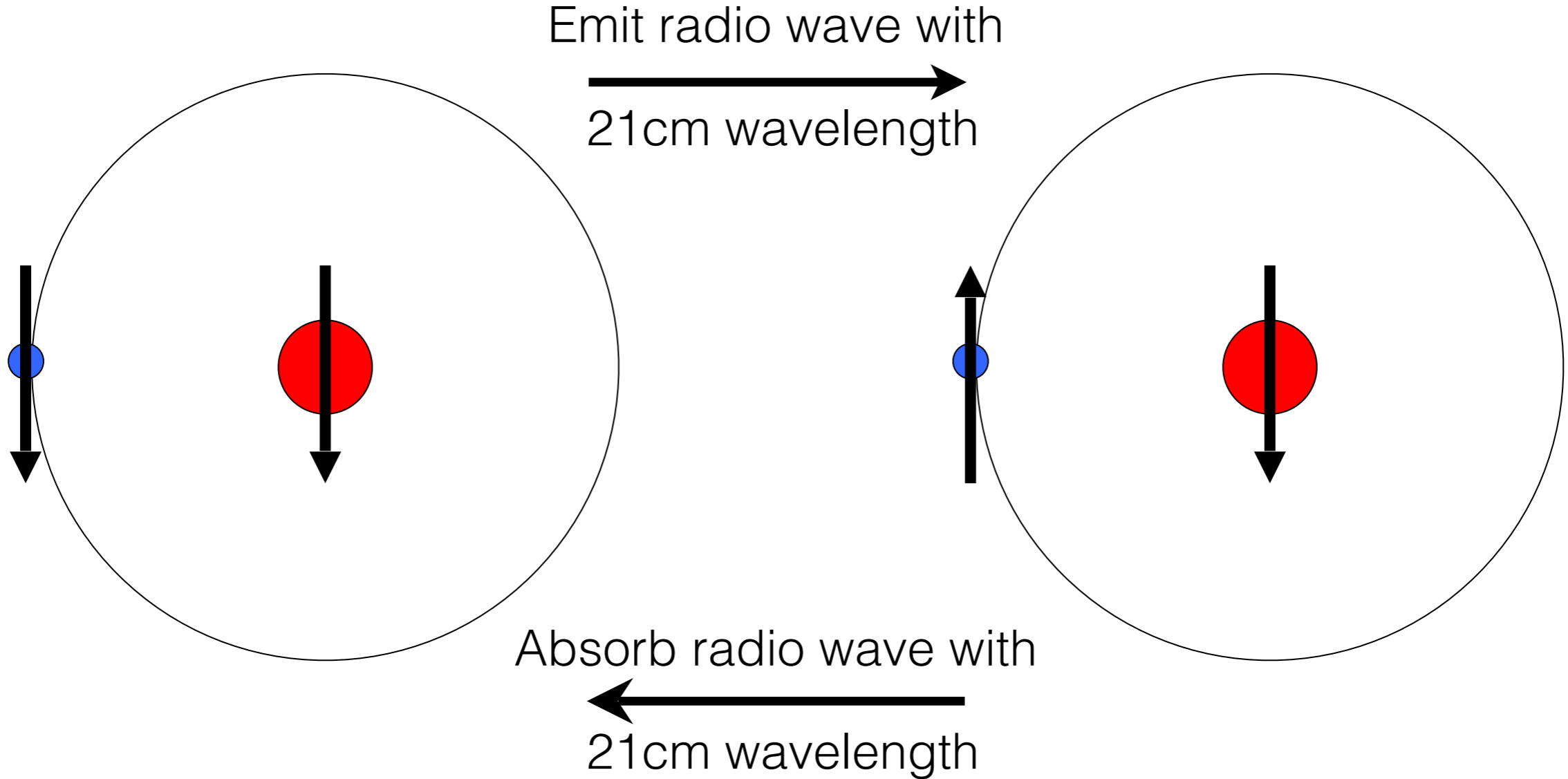
Adrian Liu  
McGill University



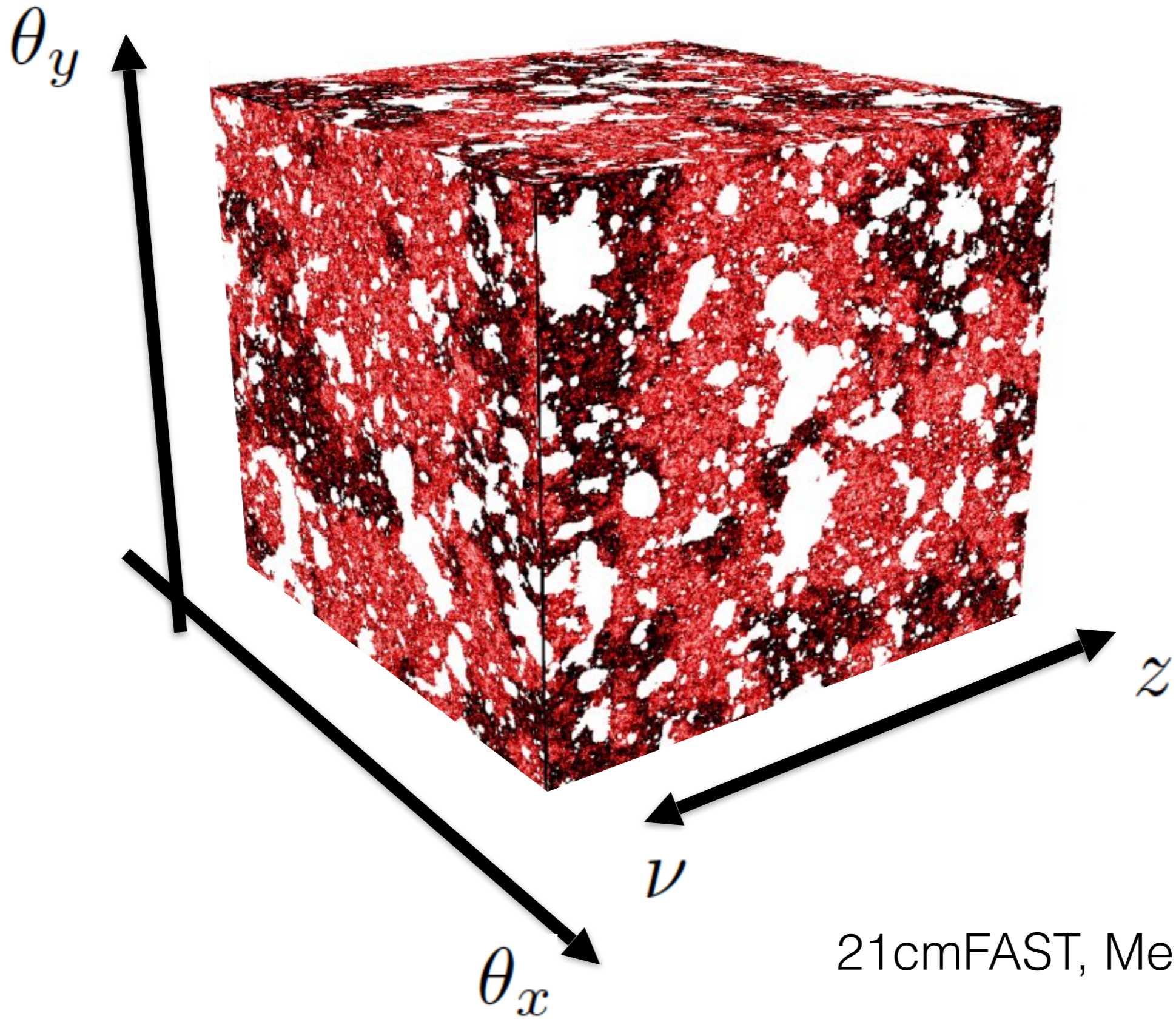
**McGill**

Vision

# Hydrogen is everywhere, and the 21cm line allows us to trace hydrogen

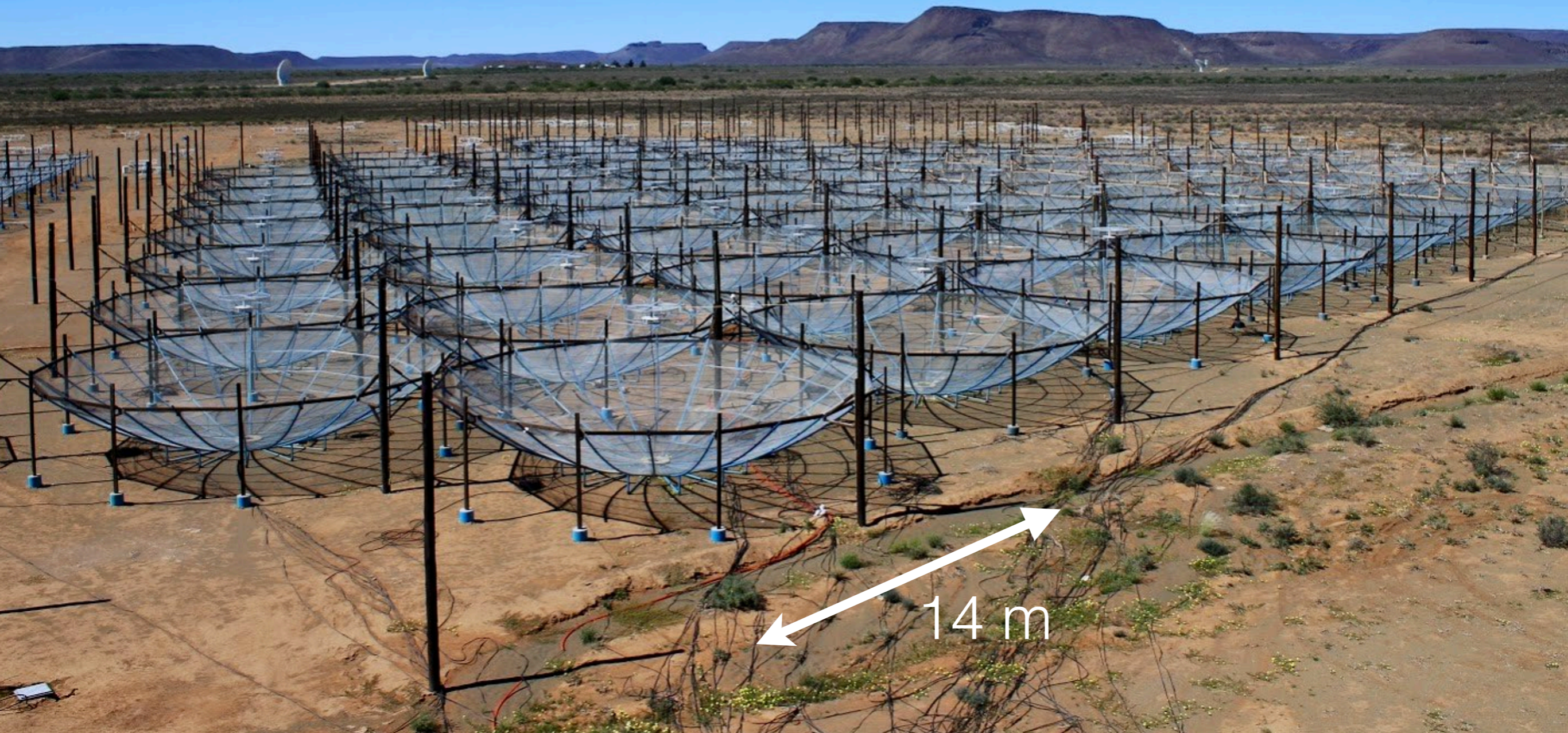


Alvarez et al. (2009)



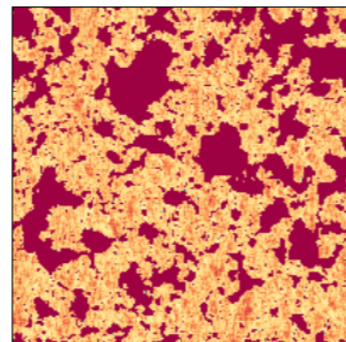
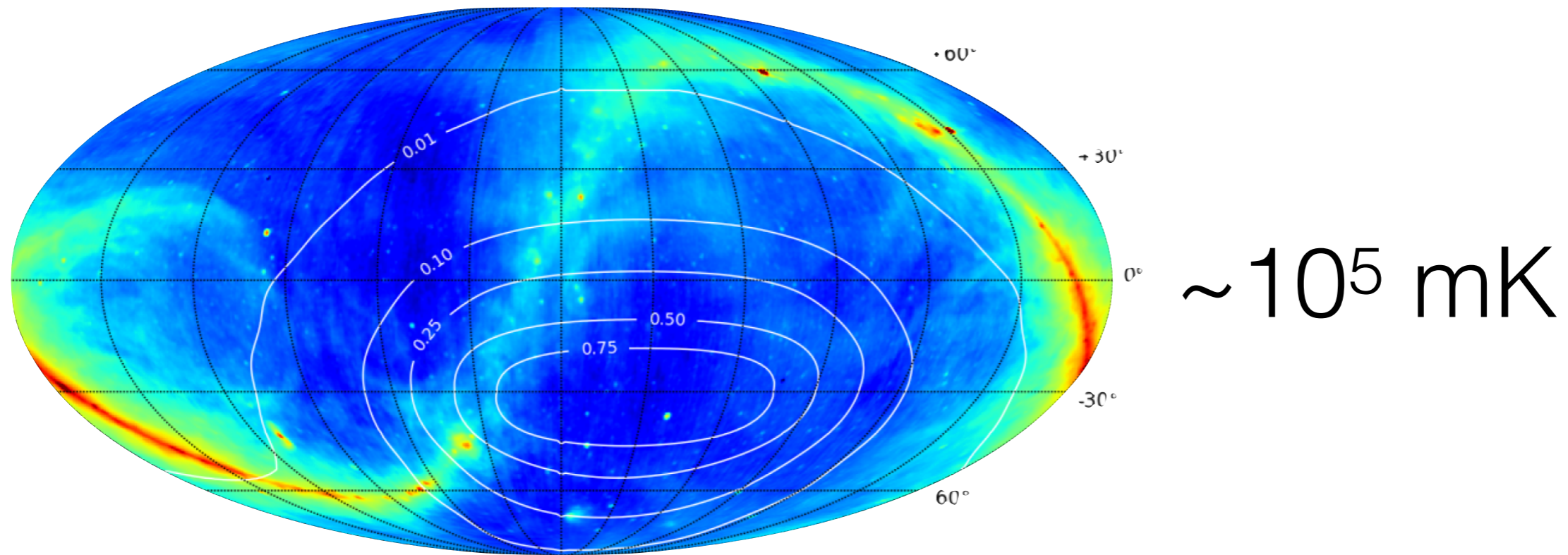
21cmFAST, Mesinger et al.

A new generation of radio observatories such as the **Hydrogen Epoch of Reionization Array (HERA)** are striving to detect these signals



The Challenge...

...is that **foreground** contaminants are **orders of magnitude brighter**

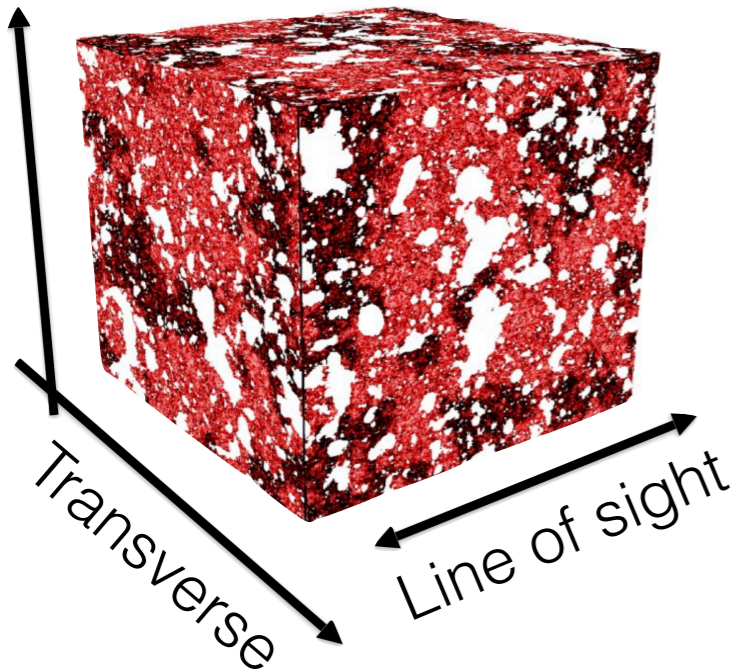


$\sim$  a few mK

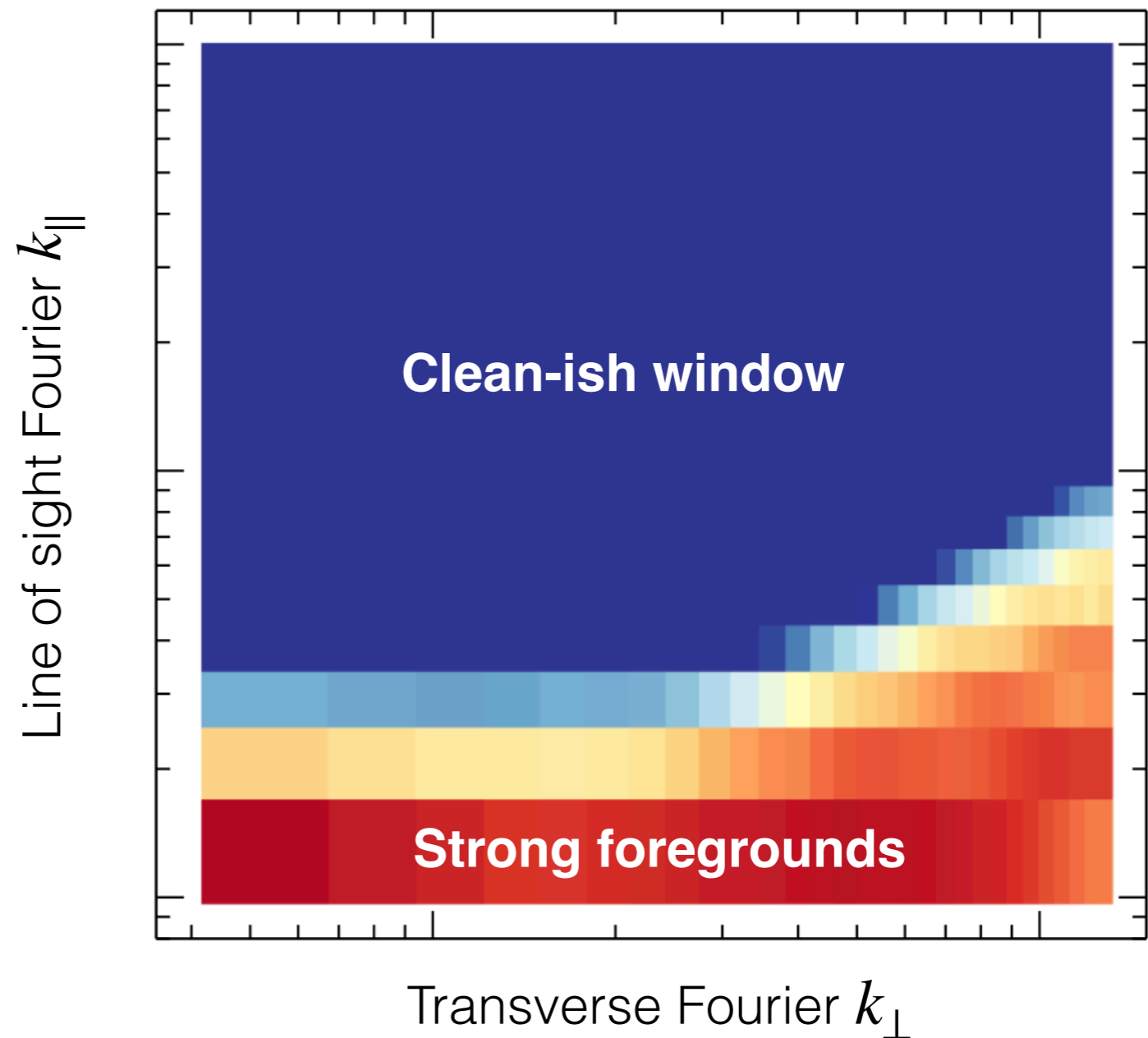
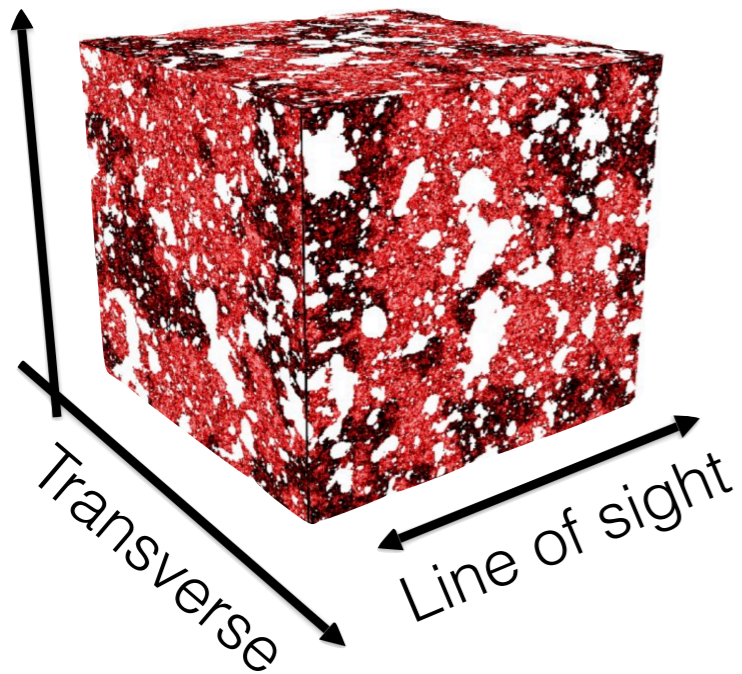


**Filtering** in Fourier space may  
be a solution...

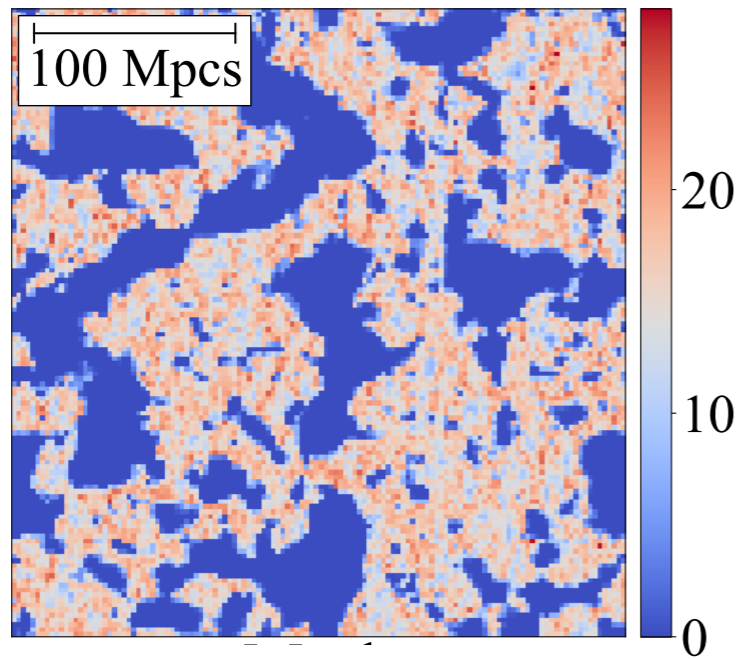
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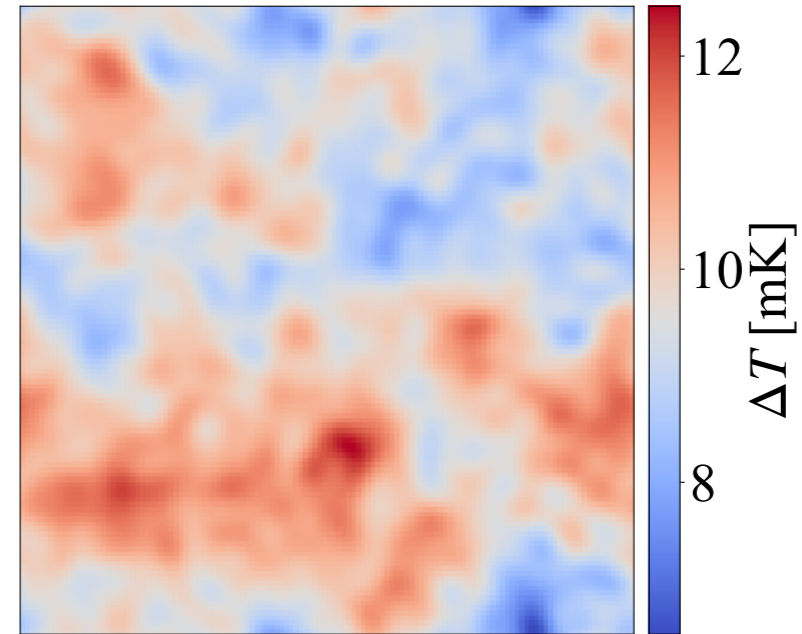
# Filtering in Fourier space may be a solution...



...but in getting rid of contaminants, it could destroy lots of information....

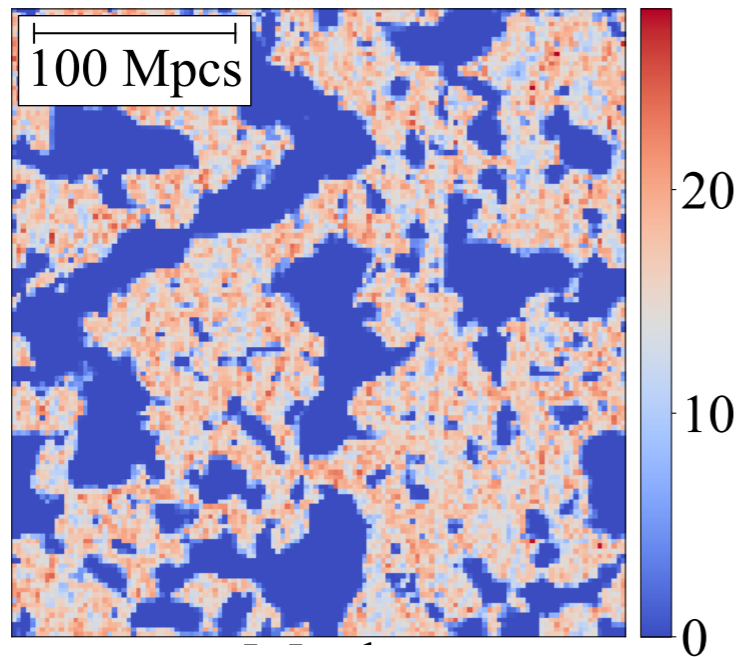


Original

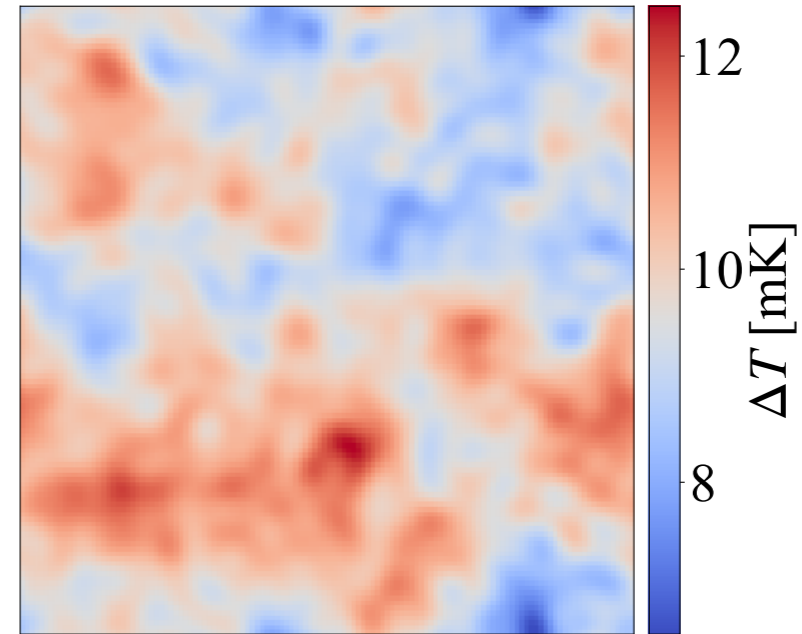


Filtered

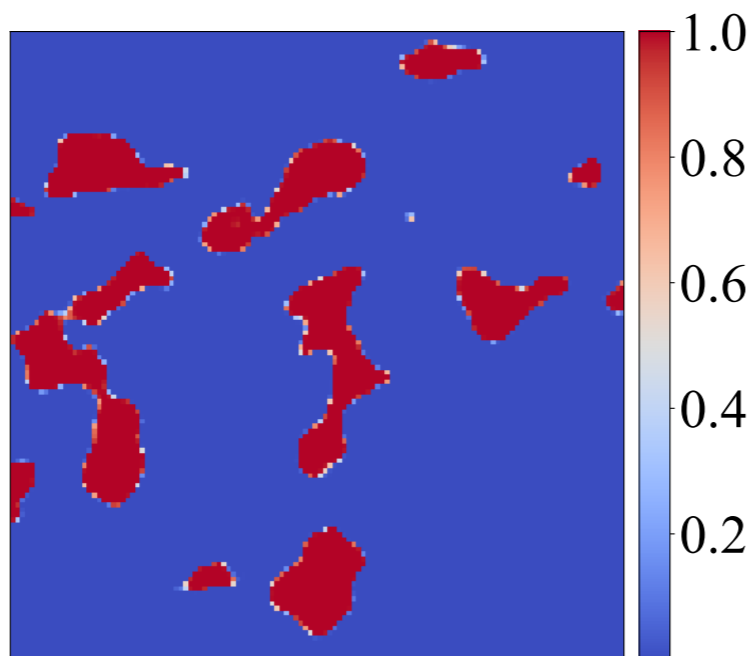
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Original



Filtered



Predicted ionization

....but perhaps machine learning and students can save the day! (Gagnon-Hartman, Cui, **AL**, Ravanbakhsh., 2021, MNRAS **504**, 4716)

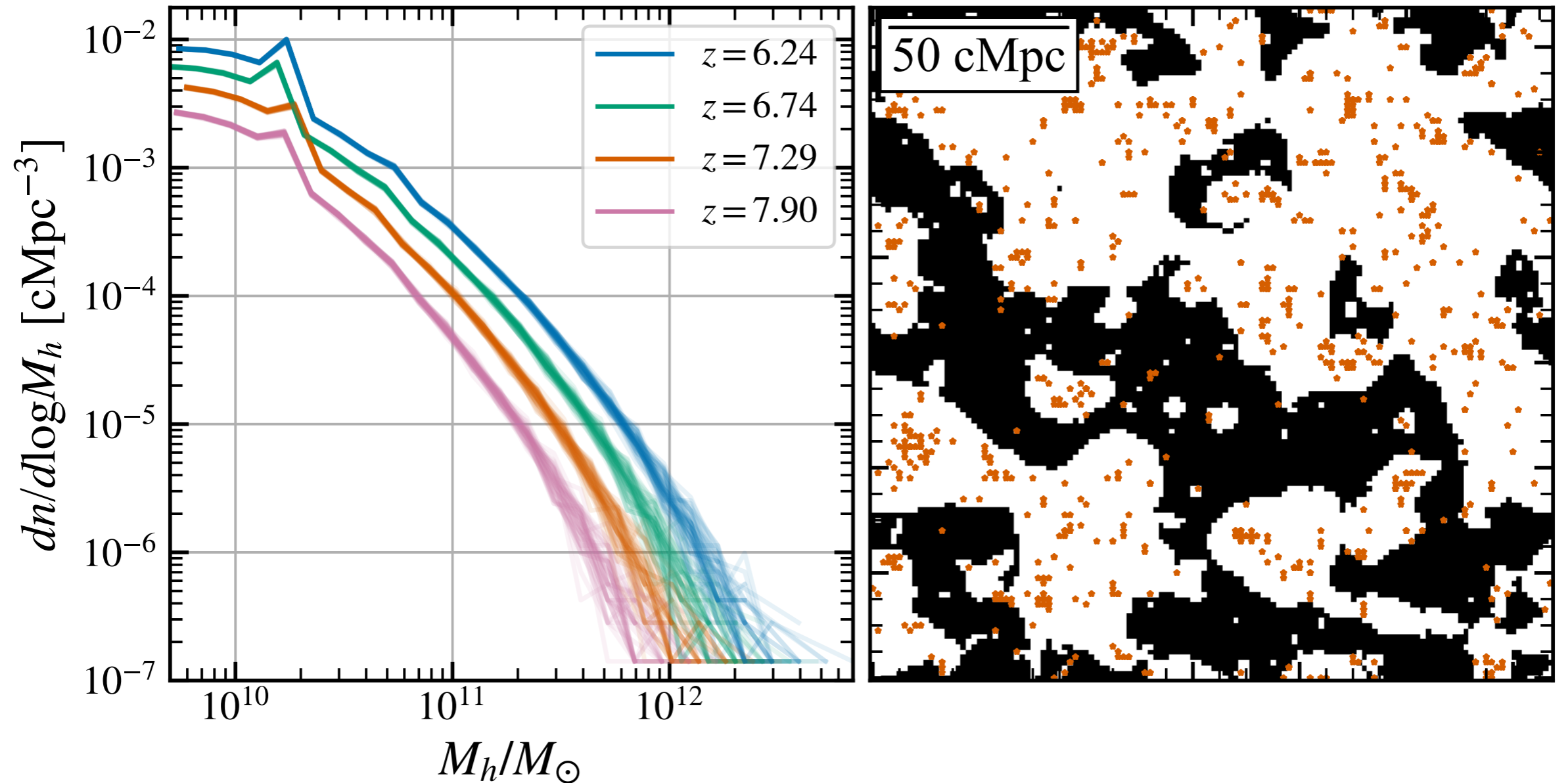


Christine Cui



Sam Gagnon-Hartman

$$\Delta T_{21}^{\text{bin}}(z = 7.29)$$



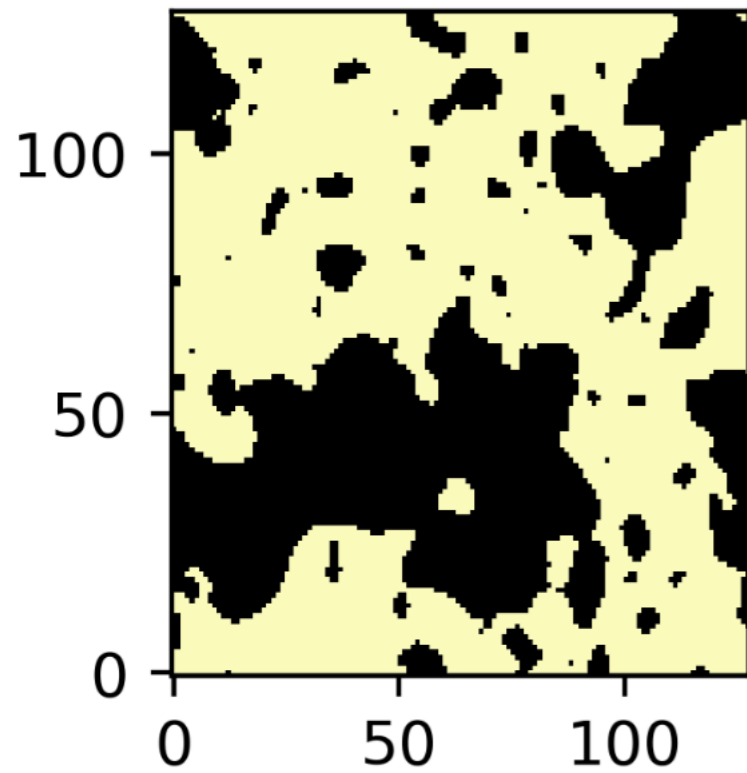
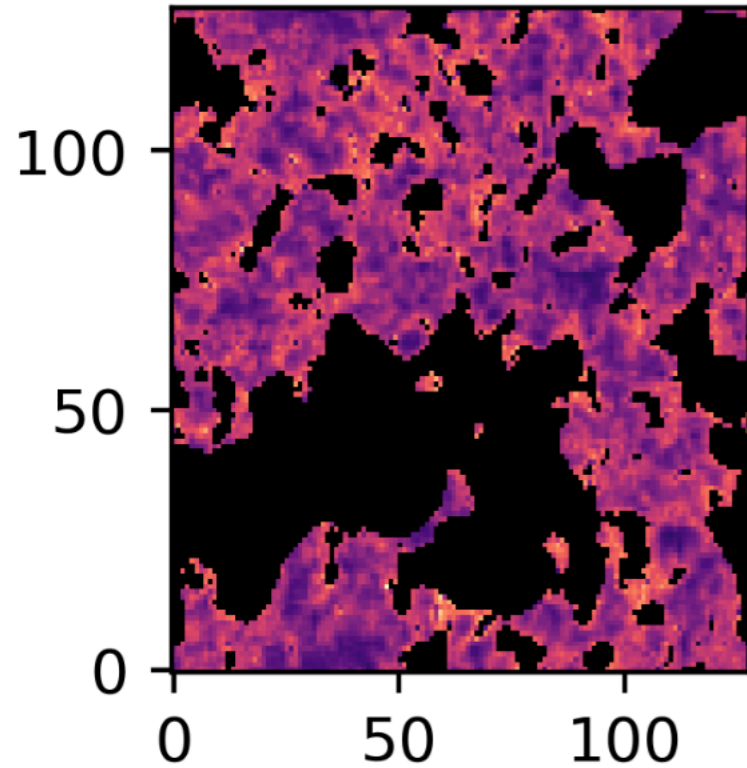
These maps can be used as  
a guide map for galaxies!

(Kennedy et al., arXiv:2308.09740)

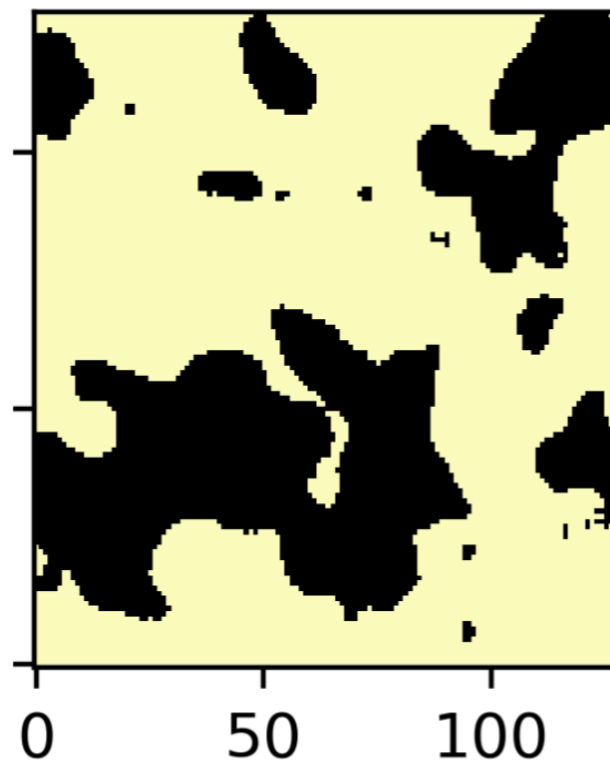


Jacob Kennedy

Or if we have already found the galaxies, we can improve the reconstruction of bubbles!



With the help of galaxies

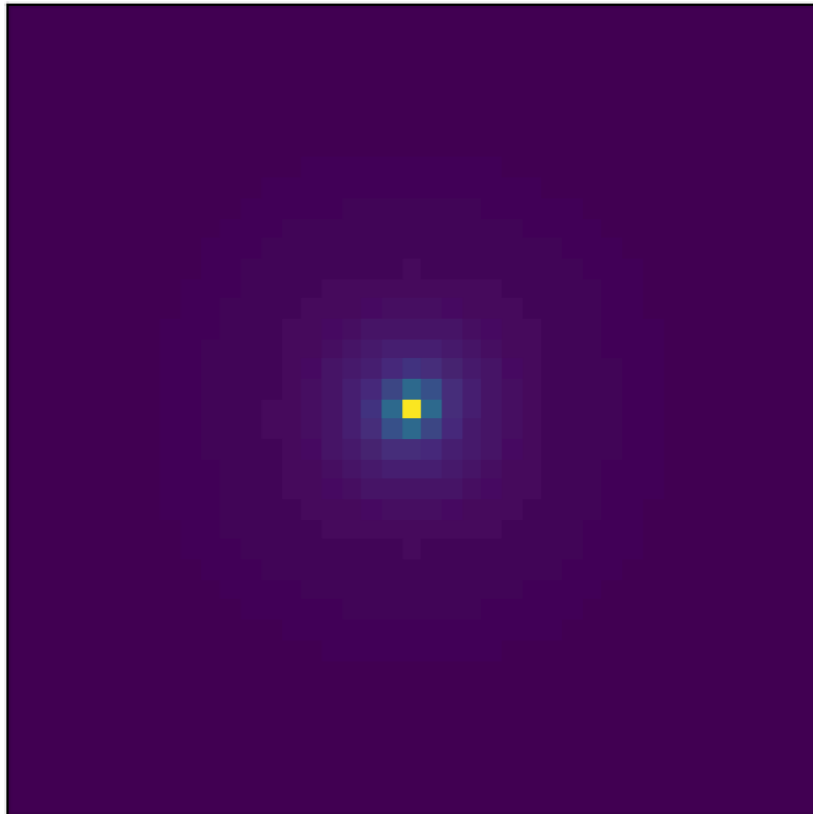


Without using galaxies

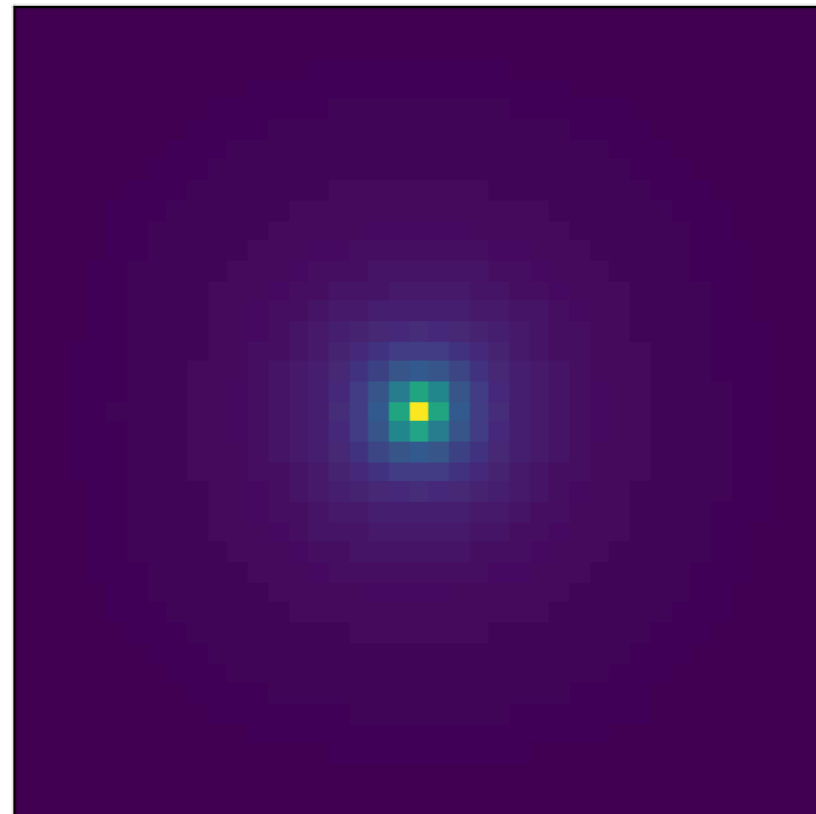


Franco del Balso

The **statistical isotropy** of our Universe means that **stacked bubbles** could be used to test the **geometry of spacetime**



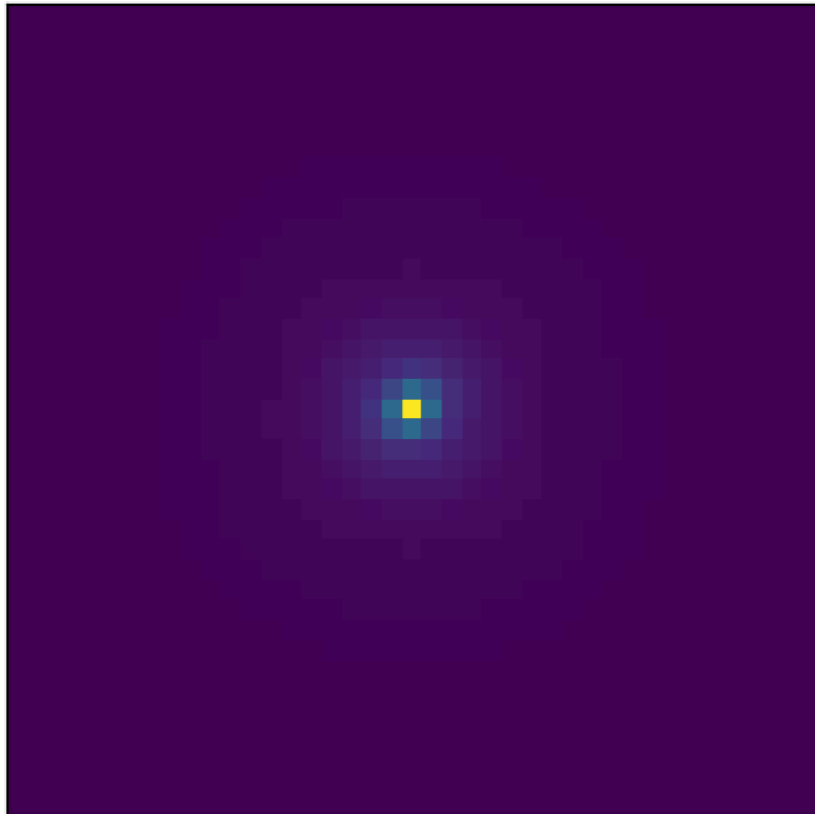
Ground truth  
bubble stack



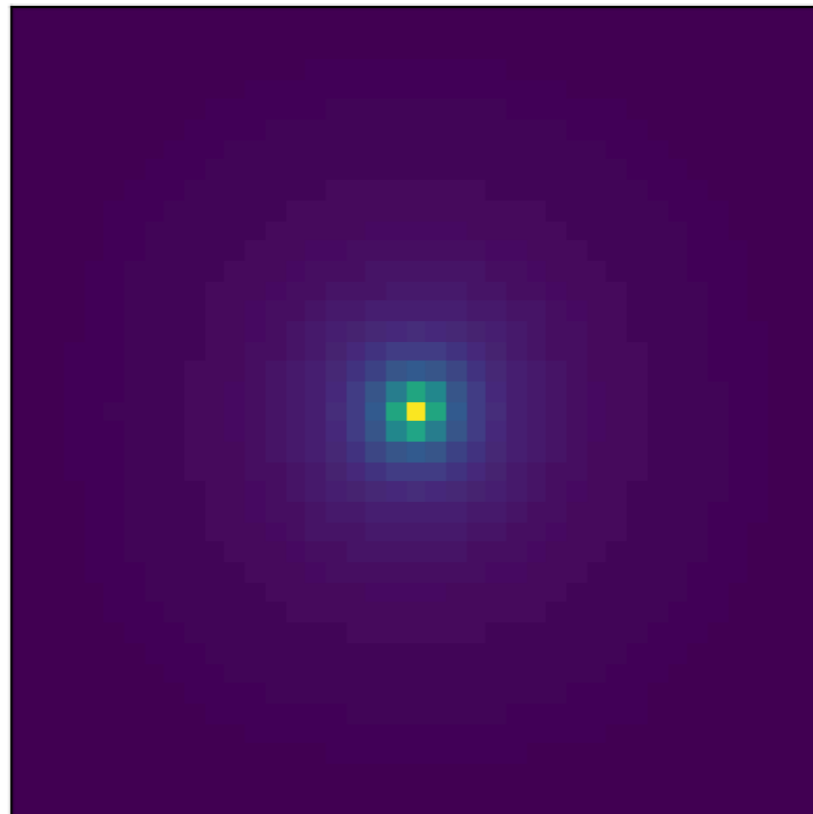
Recovered  
bubble stack



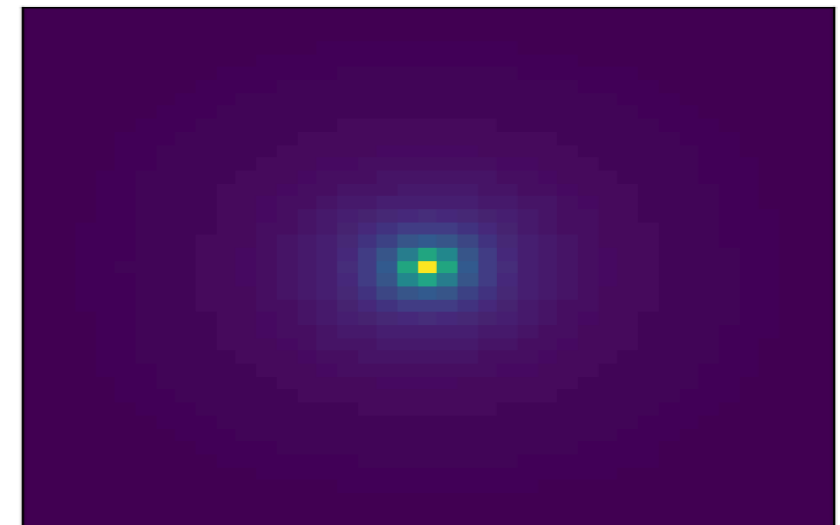
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Ground truth  
bubble stack

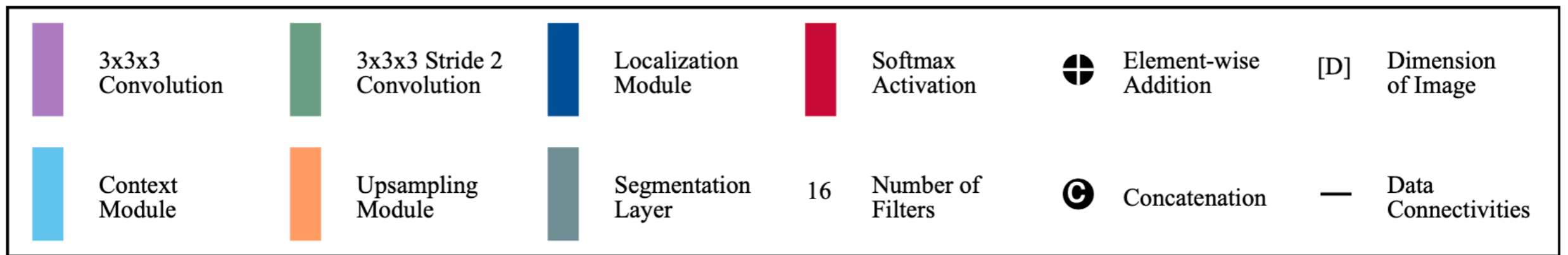
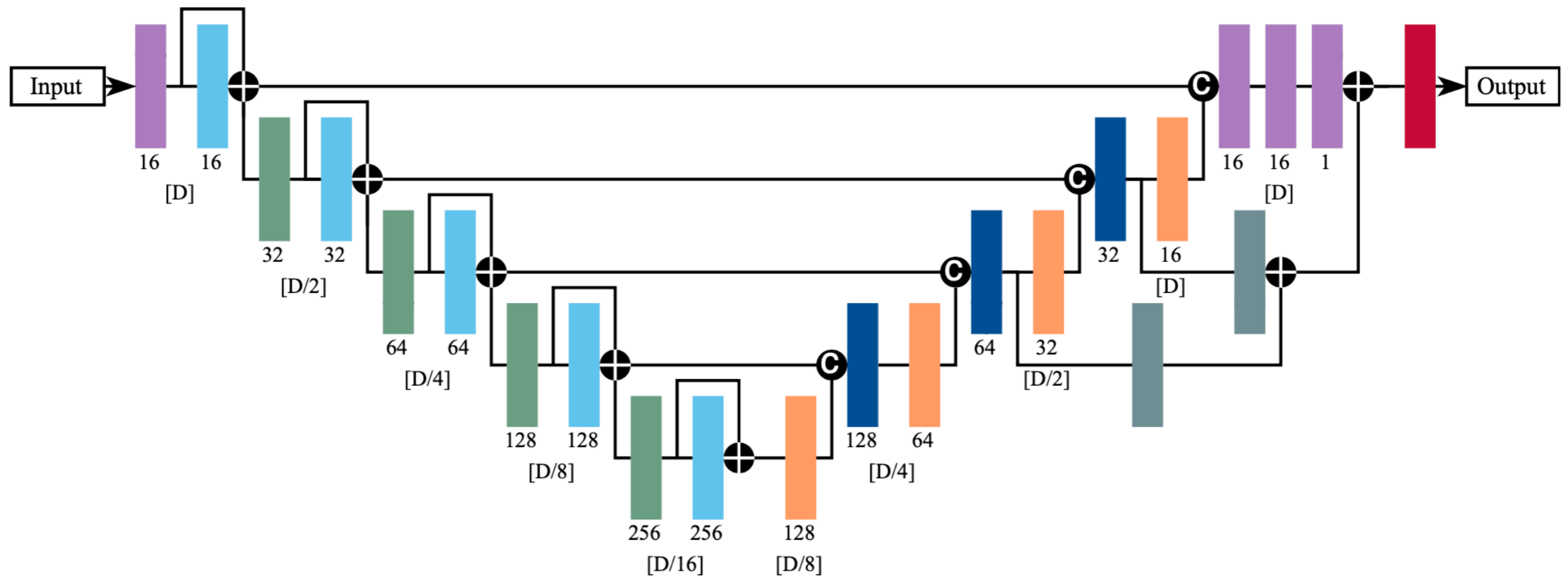


Recovered  
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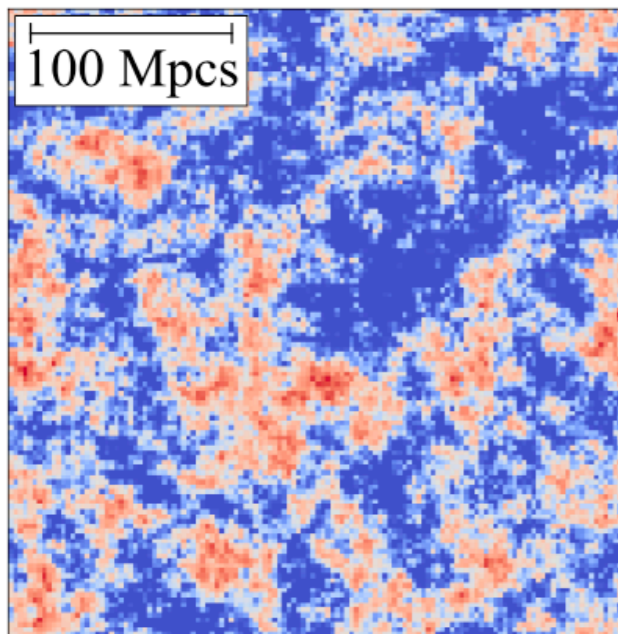


Incorrect  
assumptions about  
Universe's geometry

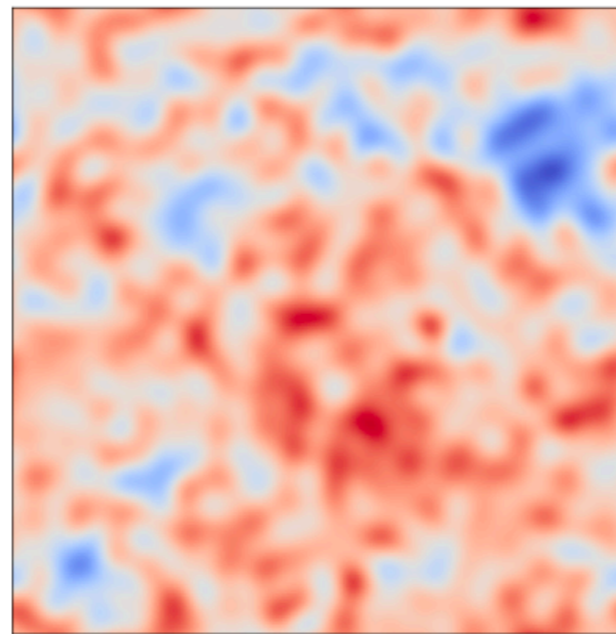
# How can this possibly even work?!



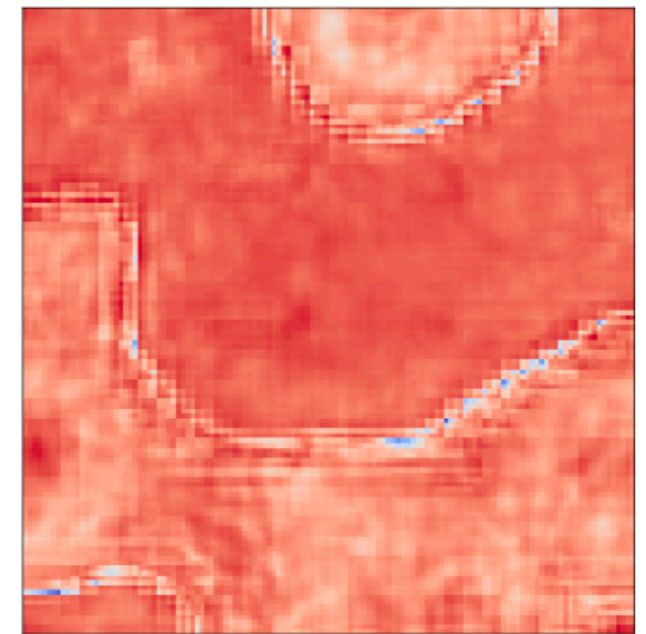
**Non-Gaussianity** to the rescue: **Gaussian** maps have **uncorrelated Fourier modes** due to stationarity, and thus **no recovery is possible**



Gaussianized  
ground truth



Filtered input



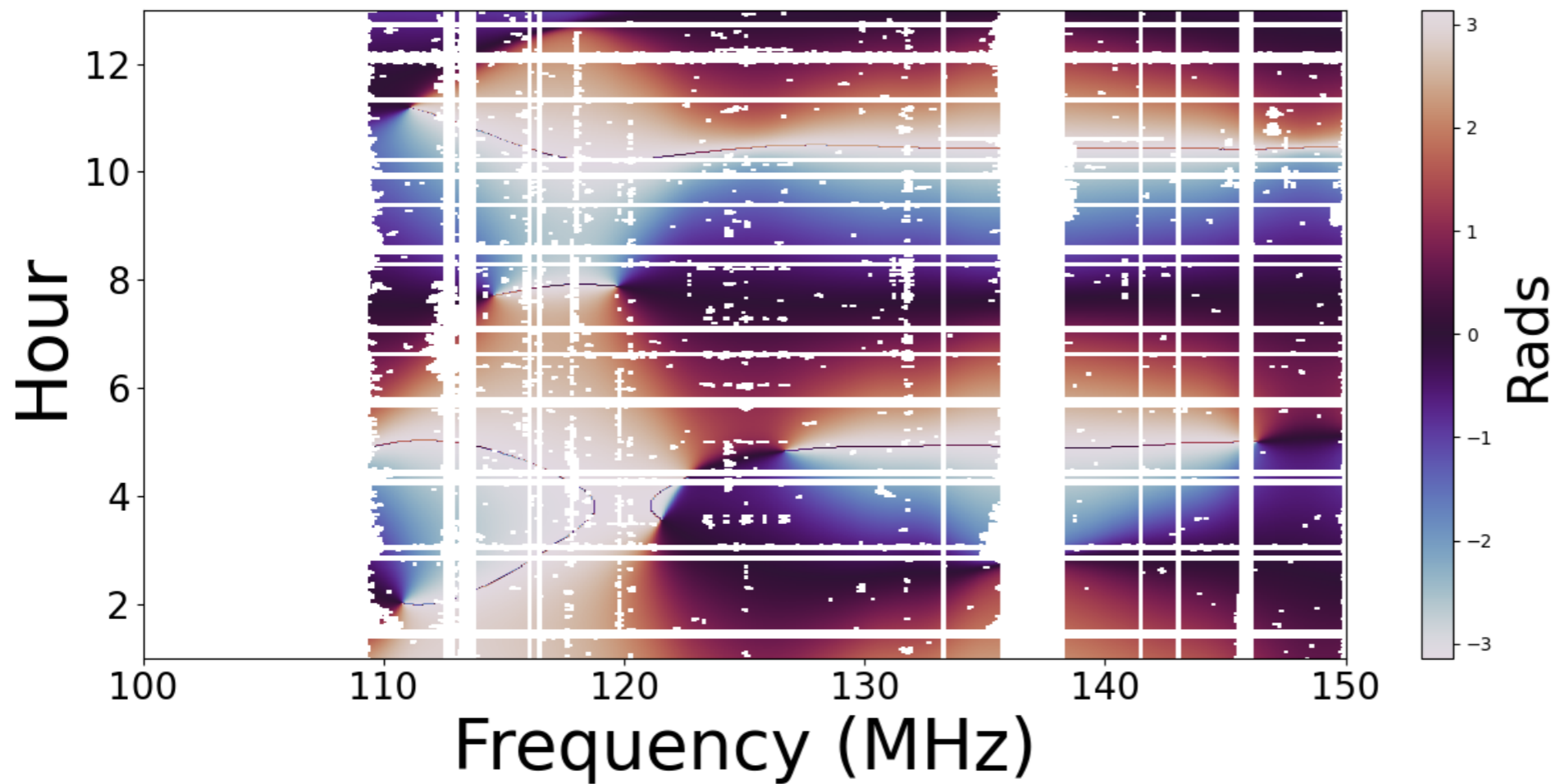
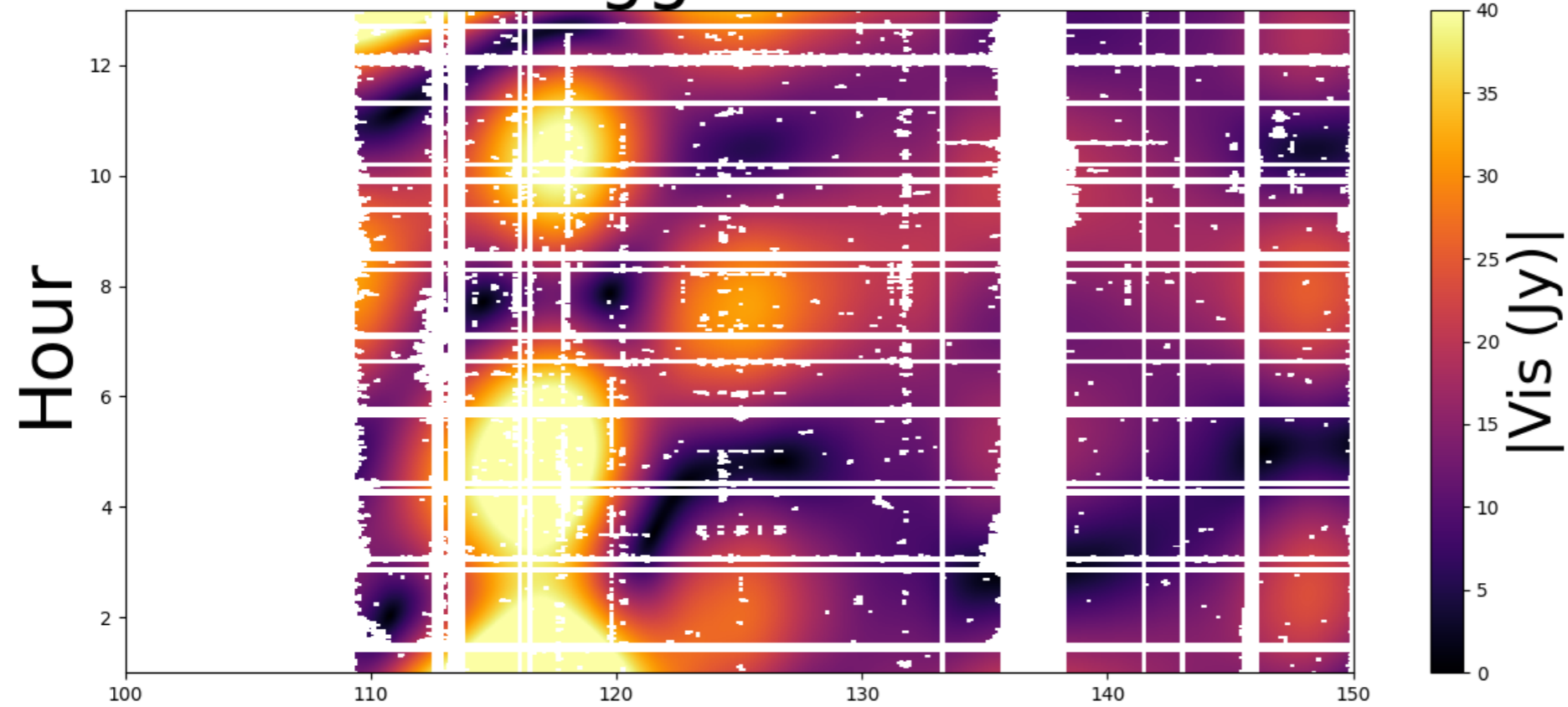
(Failed) prediction

# Other astrophysics problems in 21cm cosmology that I'd love to talk about

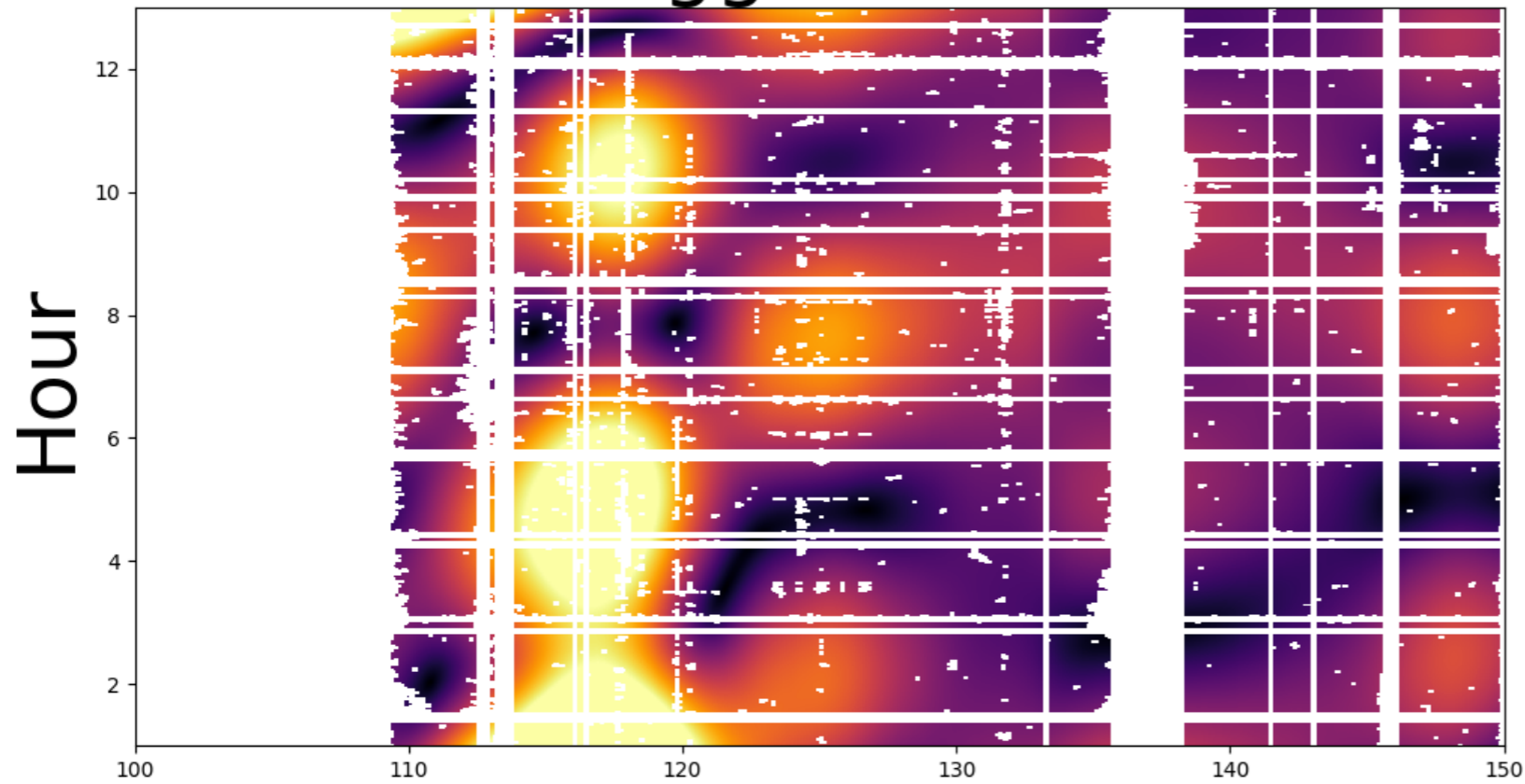
- How do we deal with the avalanche of data **without using data compression techniques that assume Gaussianity?**
- How do we best deal with **missing data?**
- How can we optimally relate “volts” in our instruments to sky signals when **our flux scales or sky models are uncertain?**
- What are the **optimal basis vectors for combining related data** from other probes?
- Can **field-level inferences (bypassing summary statistics)** be a practical path forward in this field?
- Can **symmetries in our data** be effectively utilized?
- What can **simulation-based inference** do for us?

**NML**

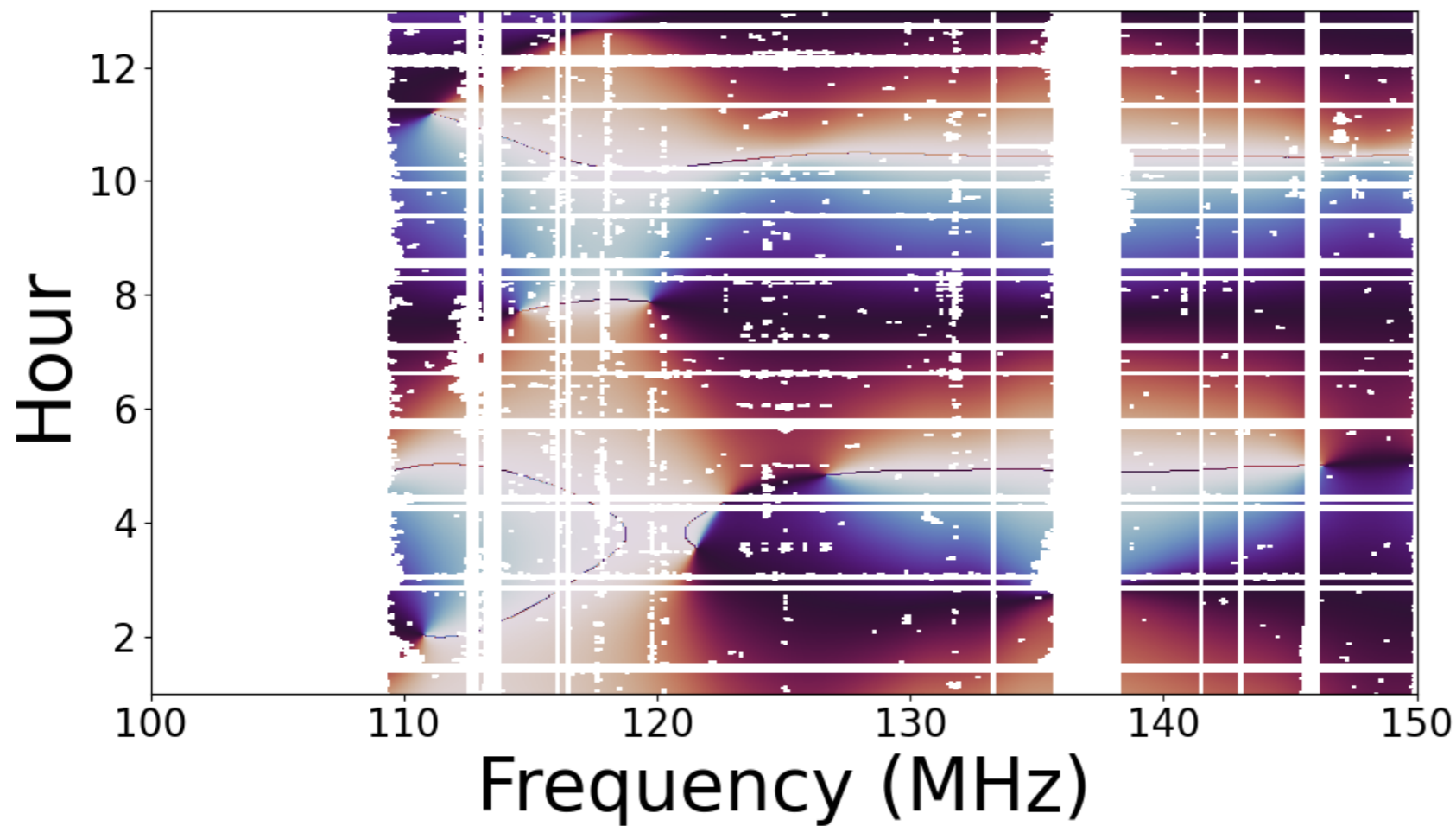
# Flagged Data



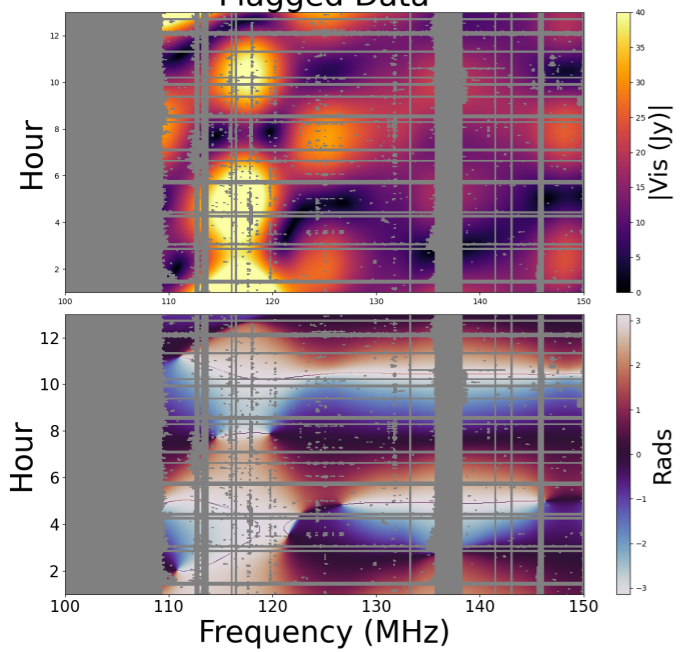
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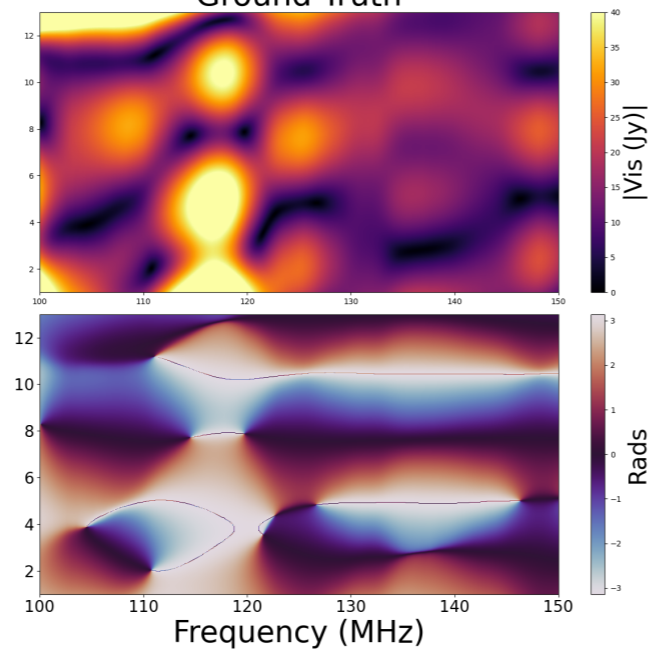
Radio  
Frequency  
Interference  
(RFI)



Flagged Data

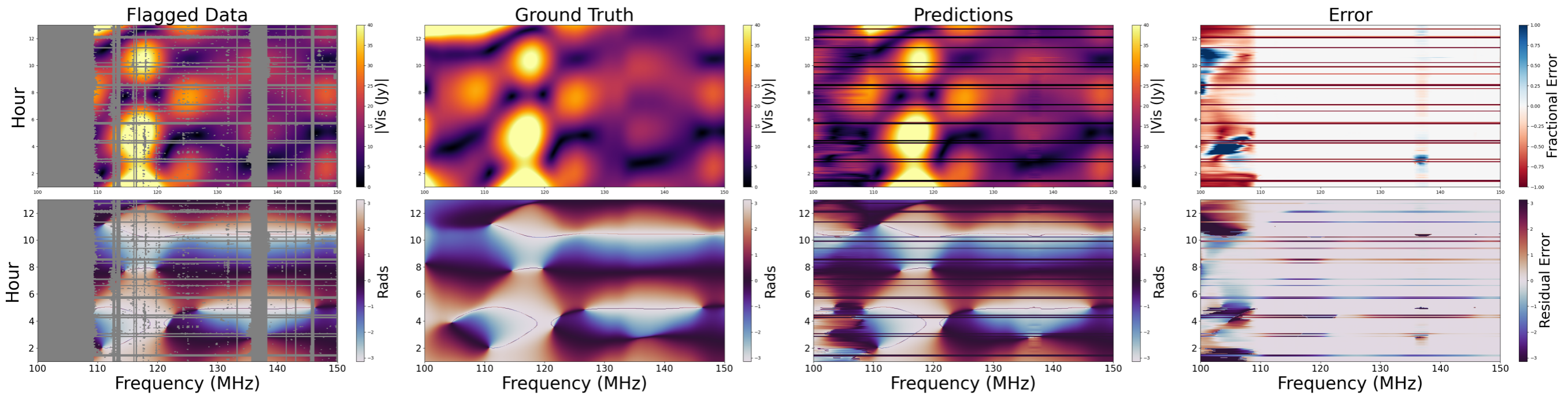


Ground Truth

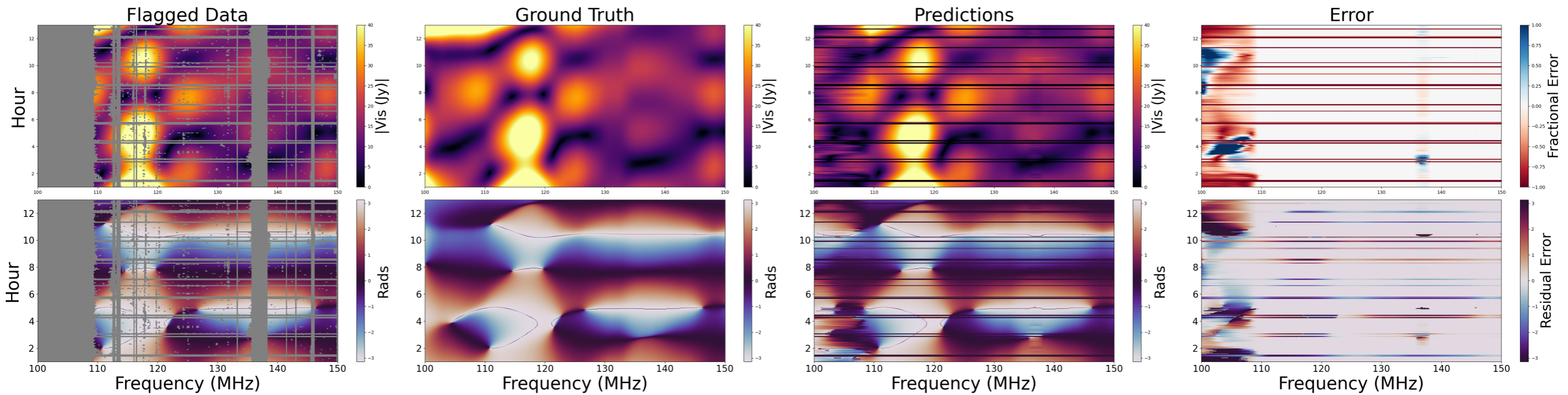




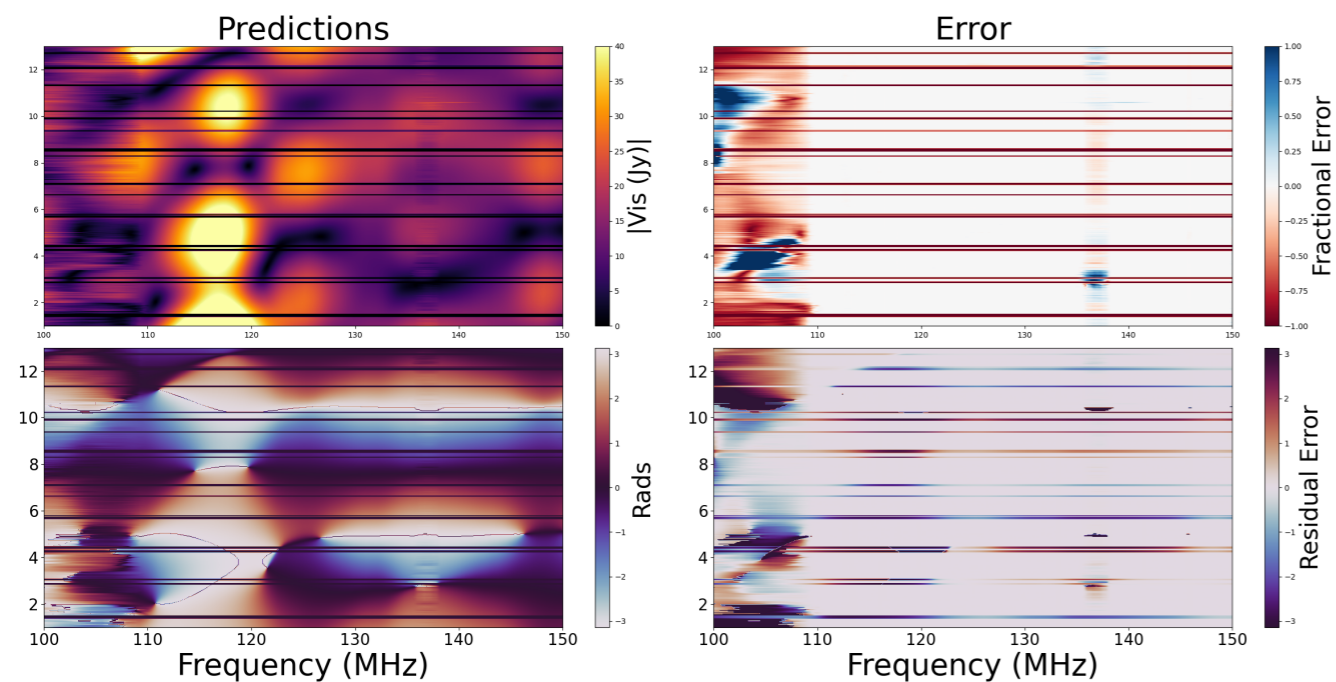
# Constrained Gaussian Realizations



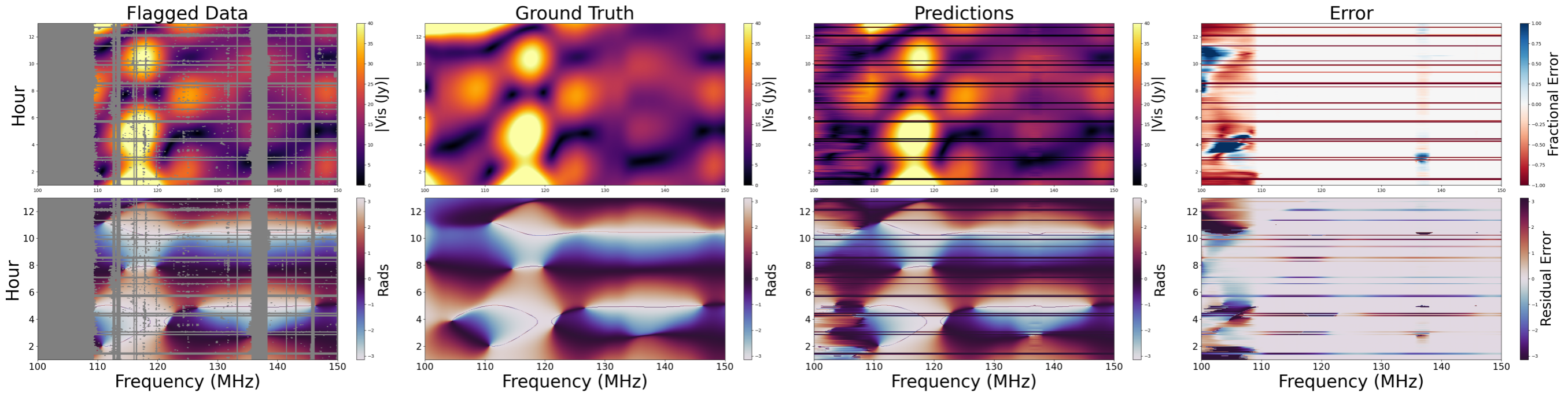
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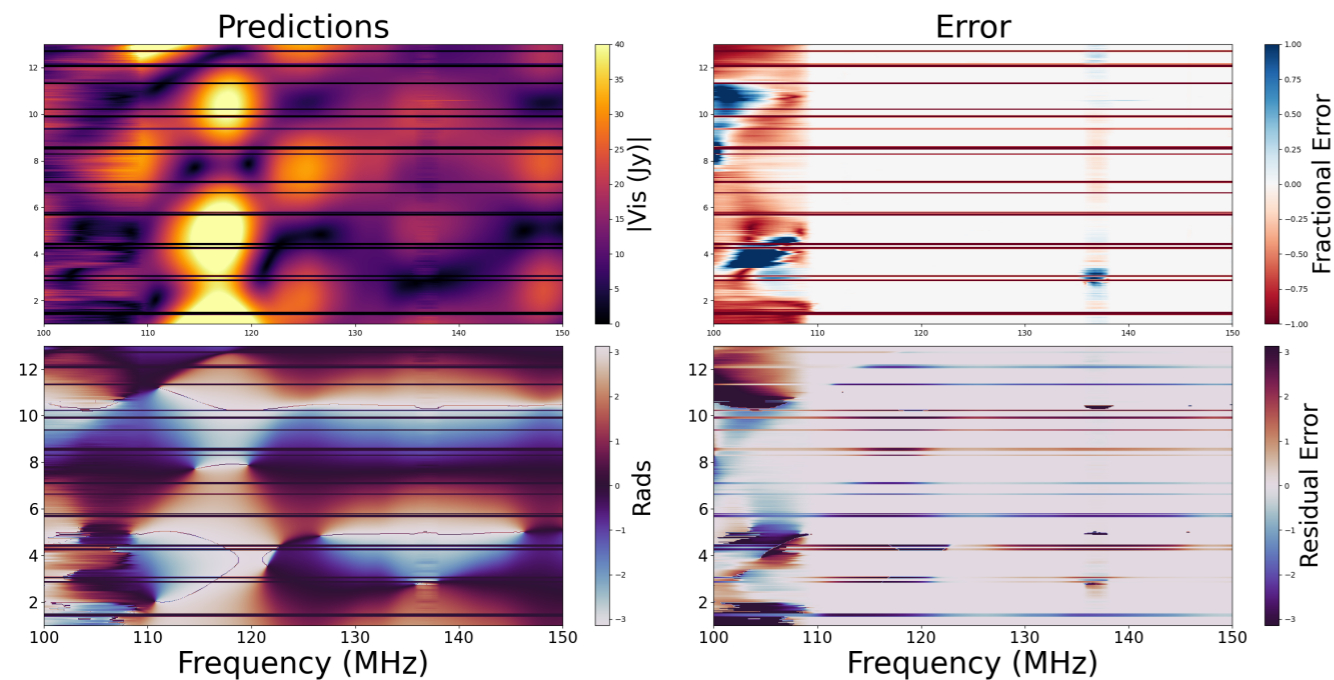
“CLEAN”



# Constrained Gaussian Realizations



“CLEAN”

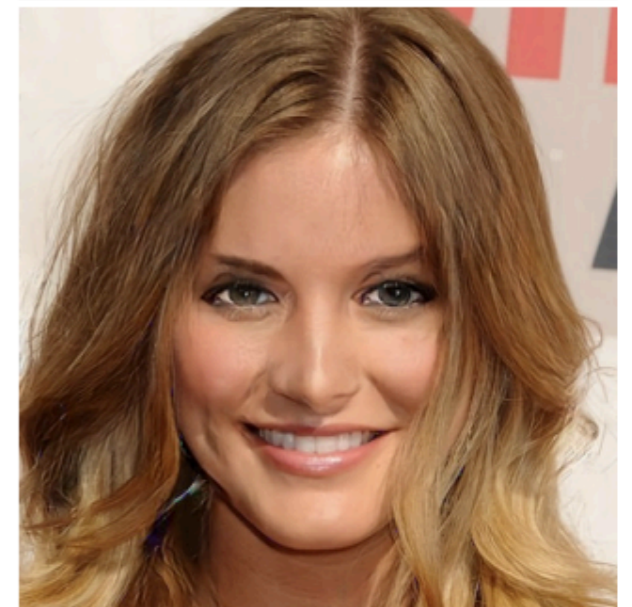
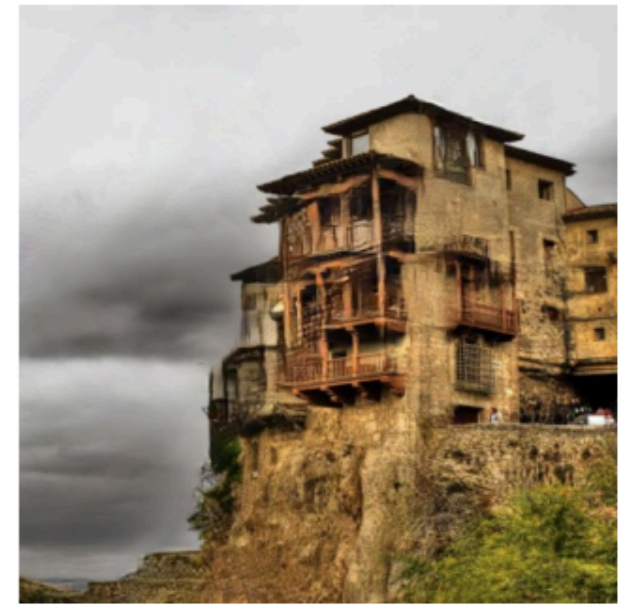


Machine Learning?

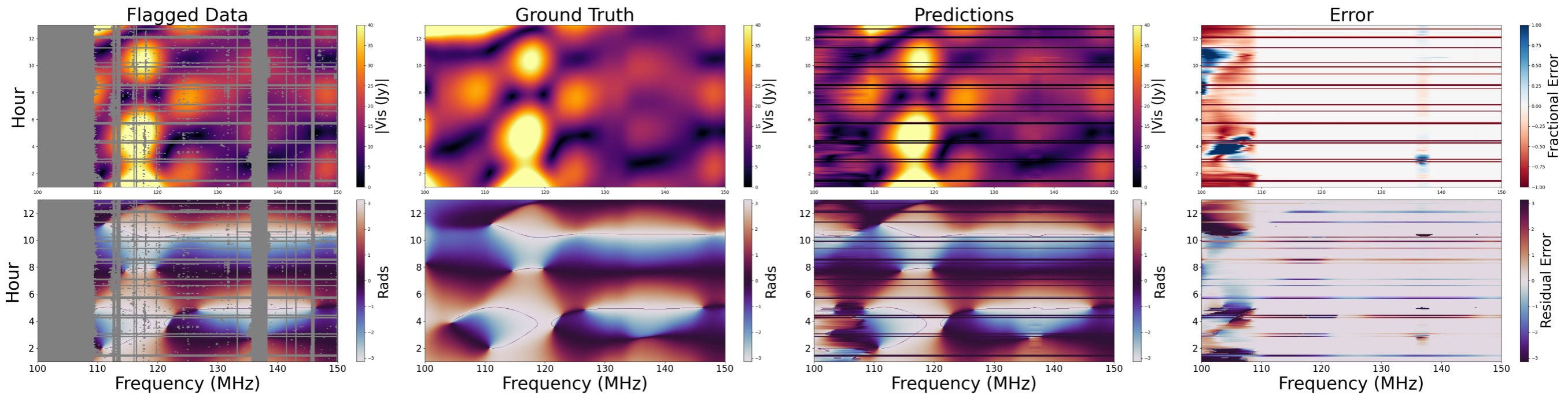


Michael  
Pagano  
PhD  
student

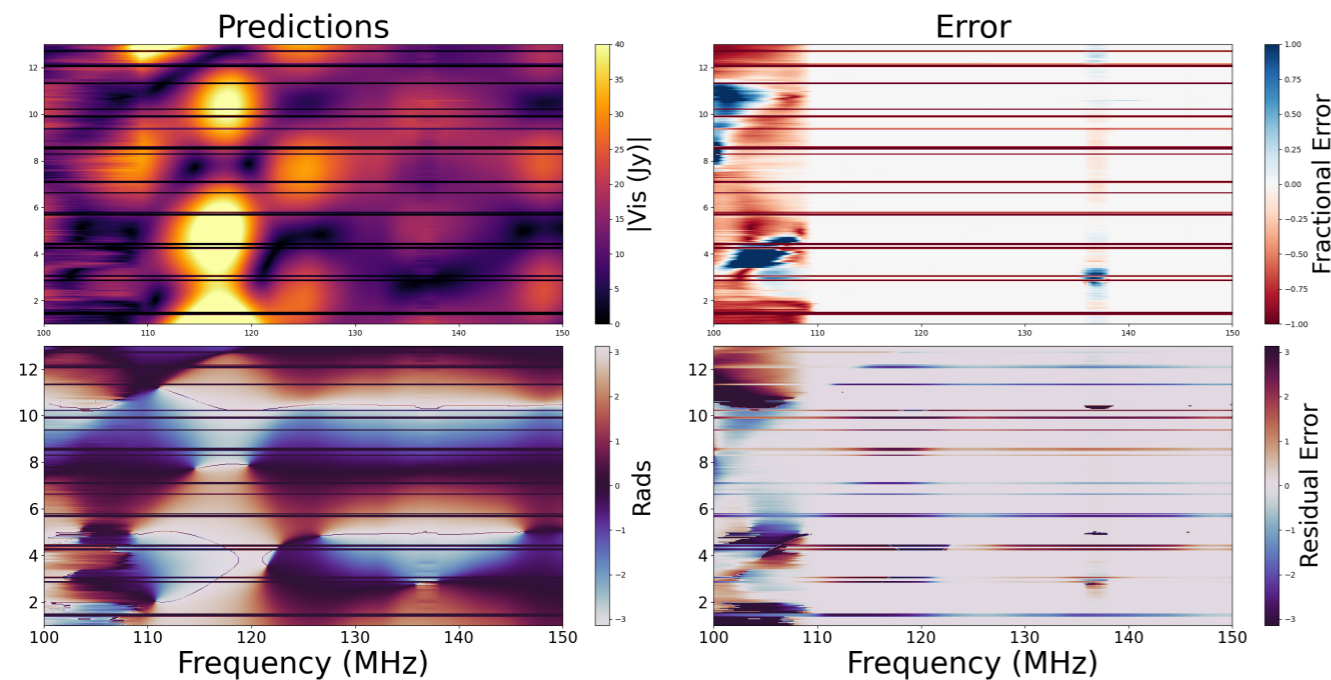
Our data is like a **picture with gaps** that can be **inpainted** using machine learning techniques



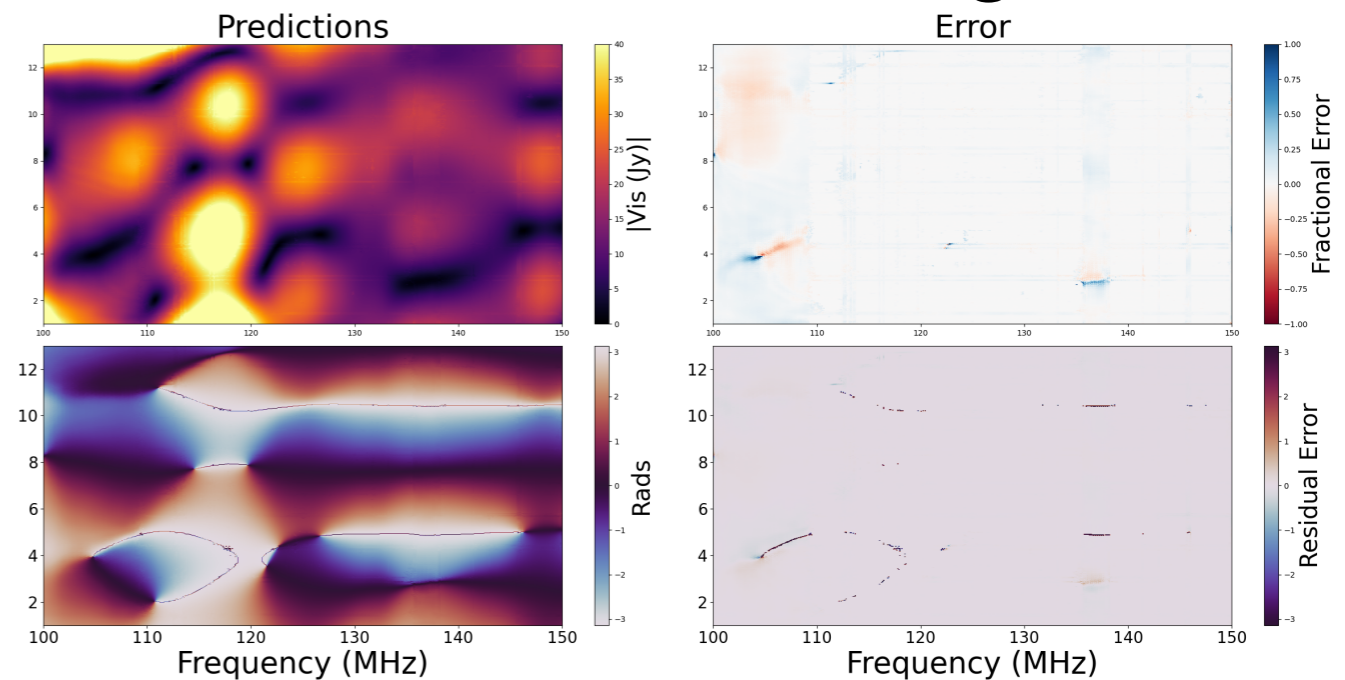
# Constrained Gaussian Realizations



“CLEAN”



Machine Learning



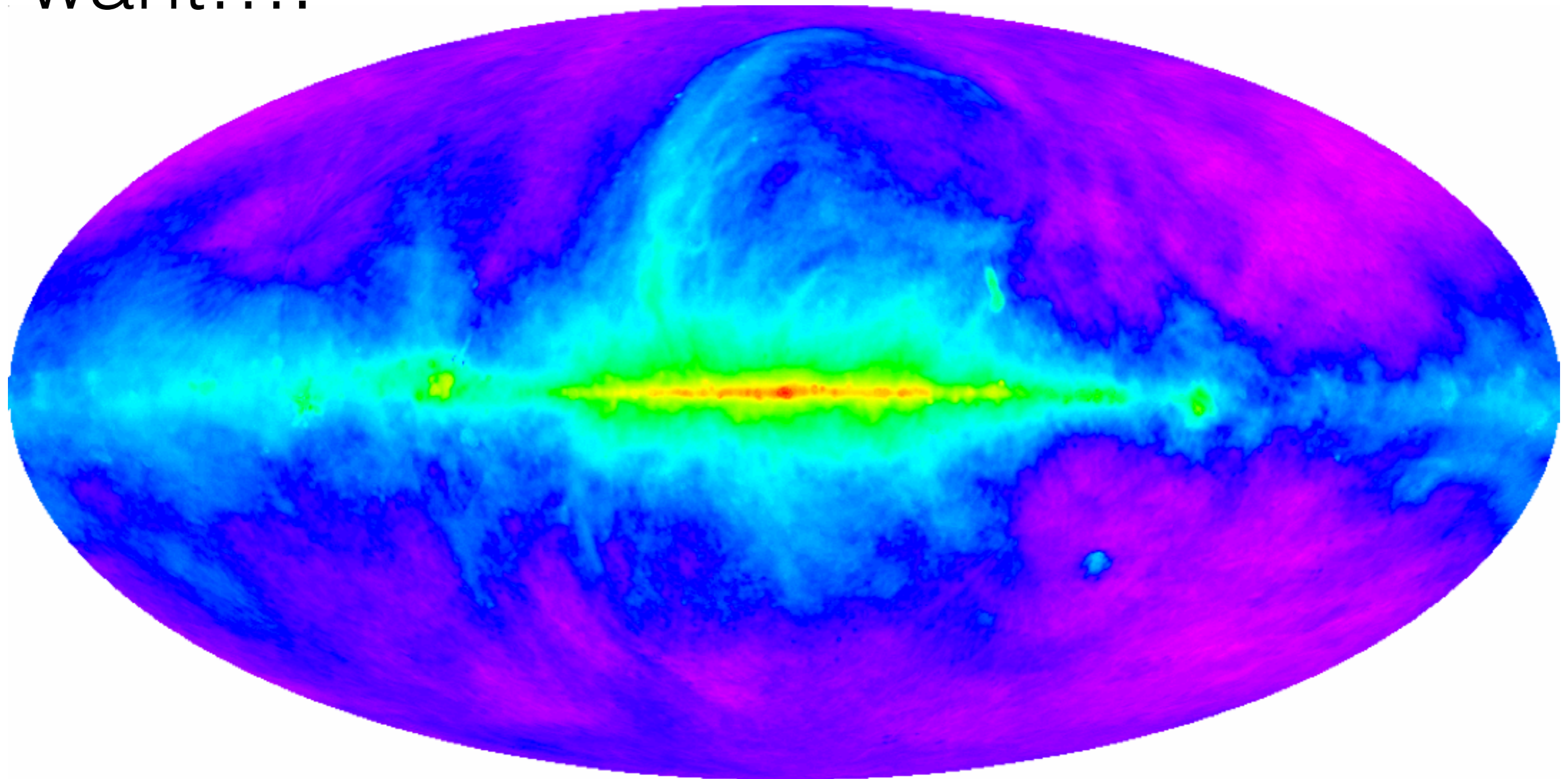
After RFI, there is an  
even greater challenge...

Our measurements are strongly  
contaminated by

**foreground emission**

that is  $\sim 10^4$  to  $10^5$  times brighter  
than the cosmological signal

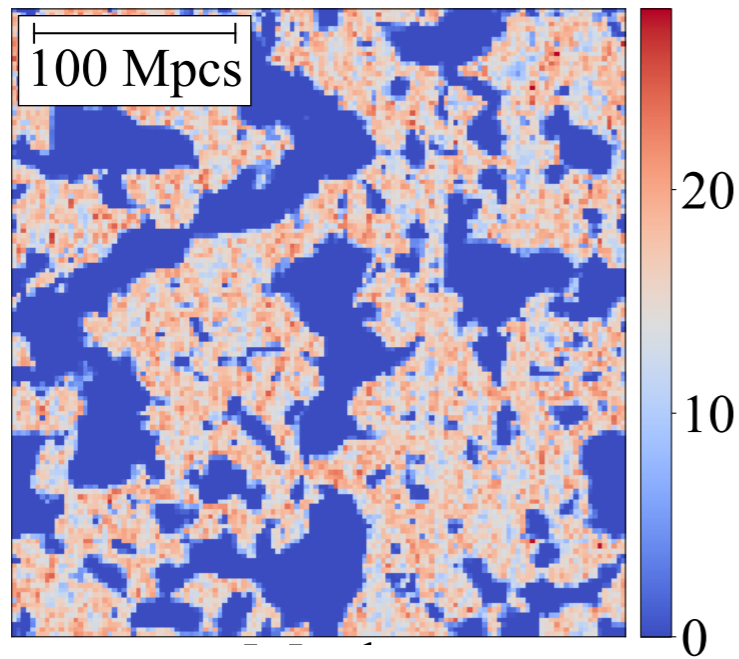
A huge portion is not what you want....



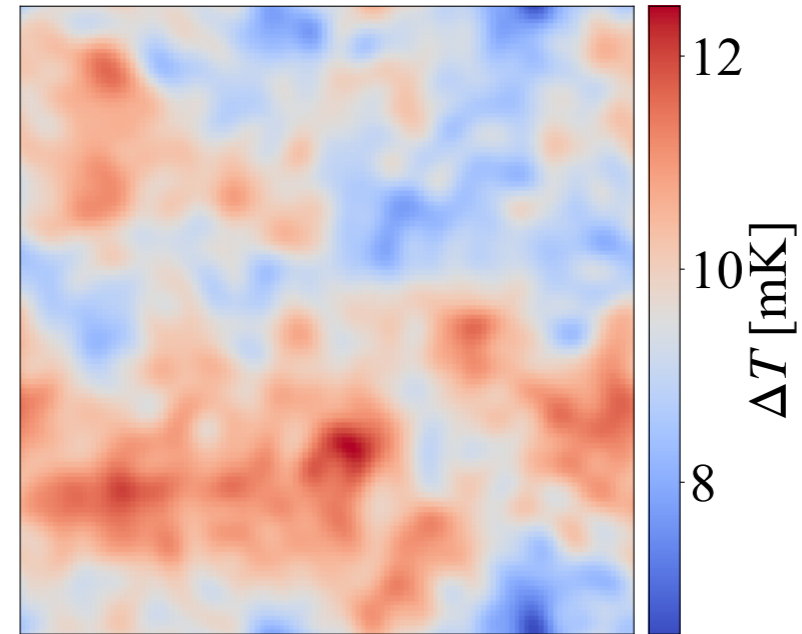
....and just a tiny portion is what you're interested in 



Filtering the data can get rid of contaminants,  
but destroy lots of information....

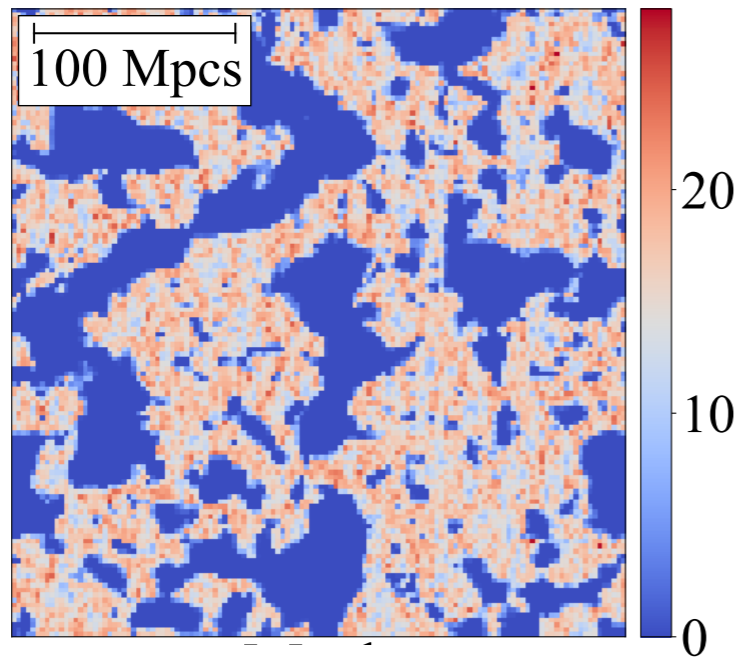


Original

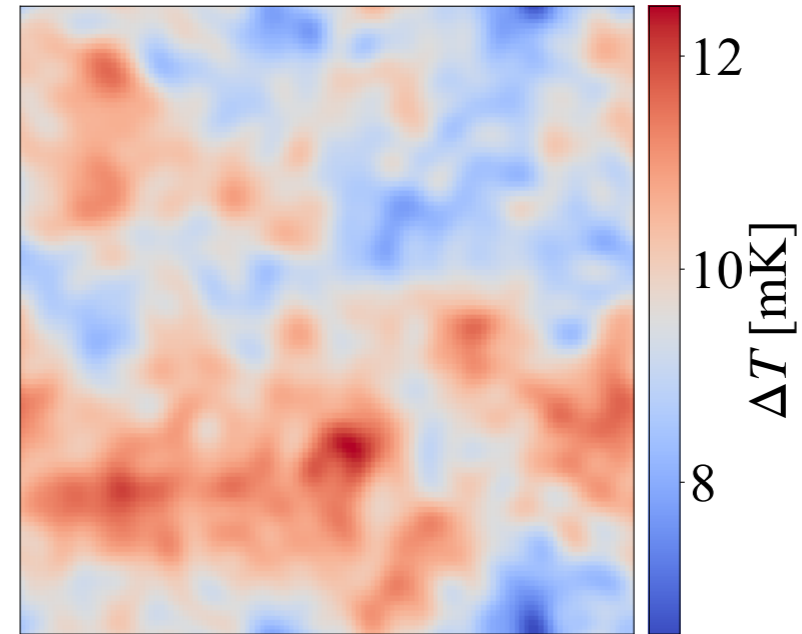


Filtered

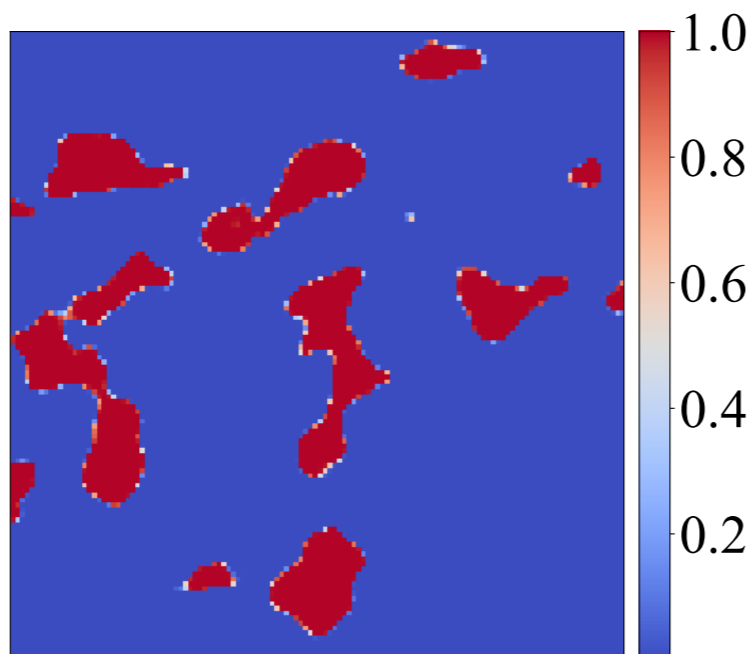
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Original



Filtered



Predicted ionization

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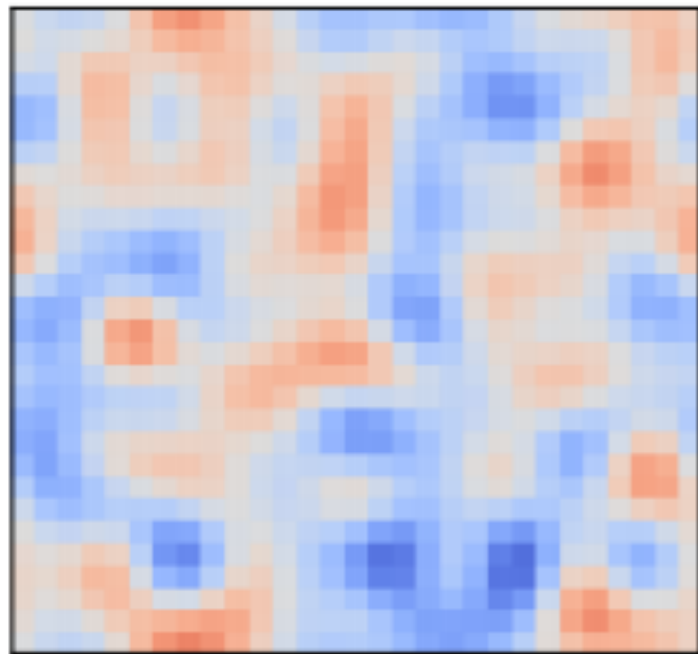


Christine Cui



Sam Gagnon-Hartman

It may even work for current-generation instruments like HERA that aren't optimized for imaging



Filtered



Original



Predicted ionization

Gagnon-Hartman, Cui, **AL**,  
Ravanbakhsh (2021)  
MNRAS **504**, 4716



Jacob Kennedy

In both the short- and the long-term,  
the best constraints will come from

**a combination of  
different probes**

Case Study #1:  
**Combining Global 21cm Signal  
Measurements with CMB Kinetic  
Sunyaev-Zel'dovich Effect  
Measurements**

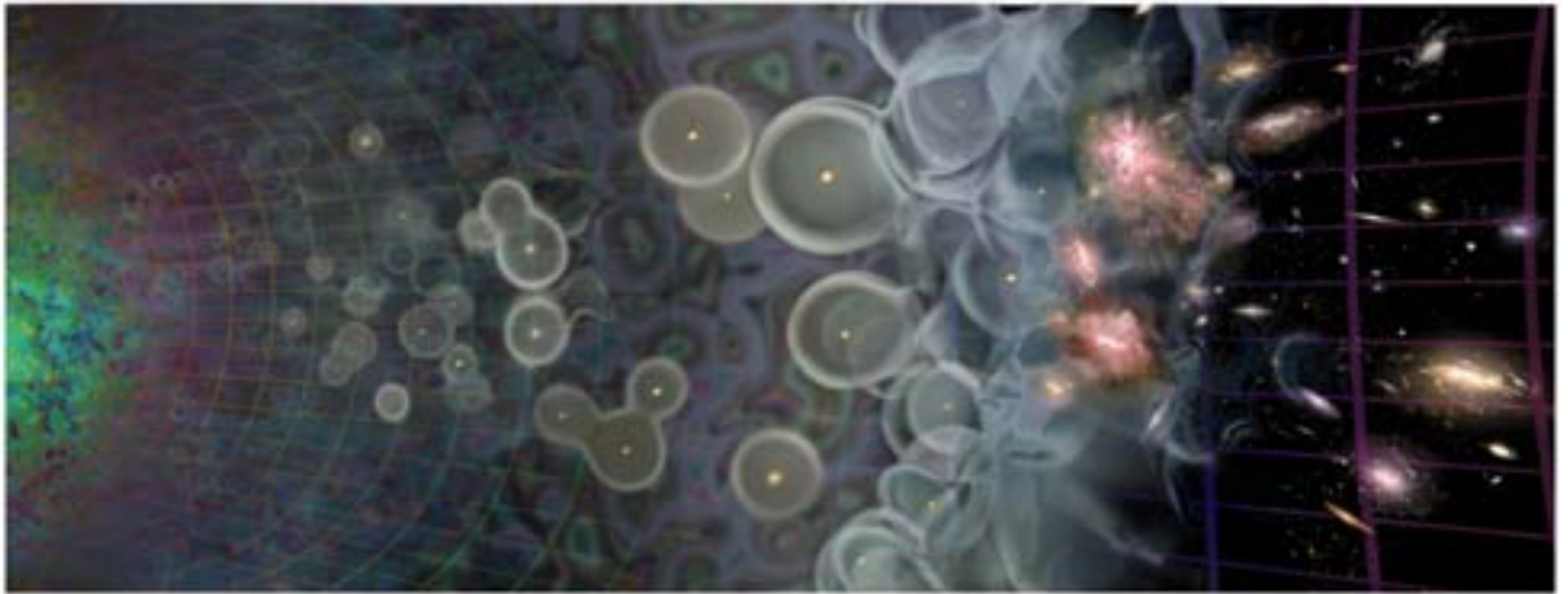


Joëlle-Marie  
Bégin

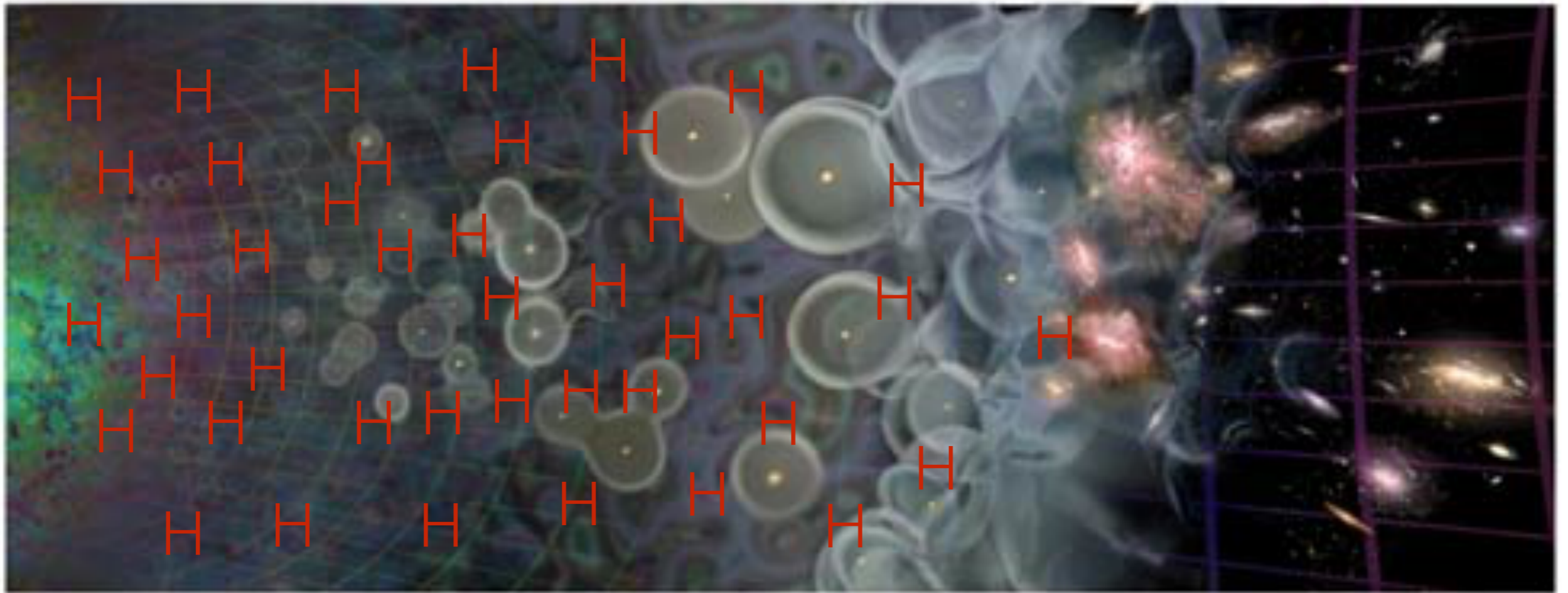
Bégin, **AL**, Gorce (2022),  
PRD **105**, 083503



Adélie  
Gorce



← Epoch of Reionization →

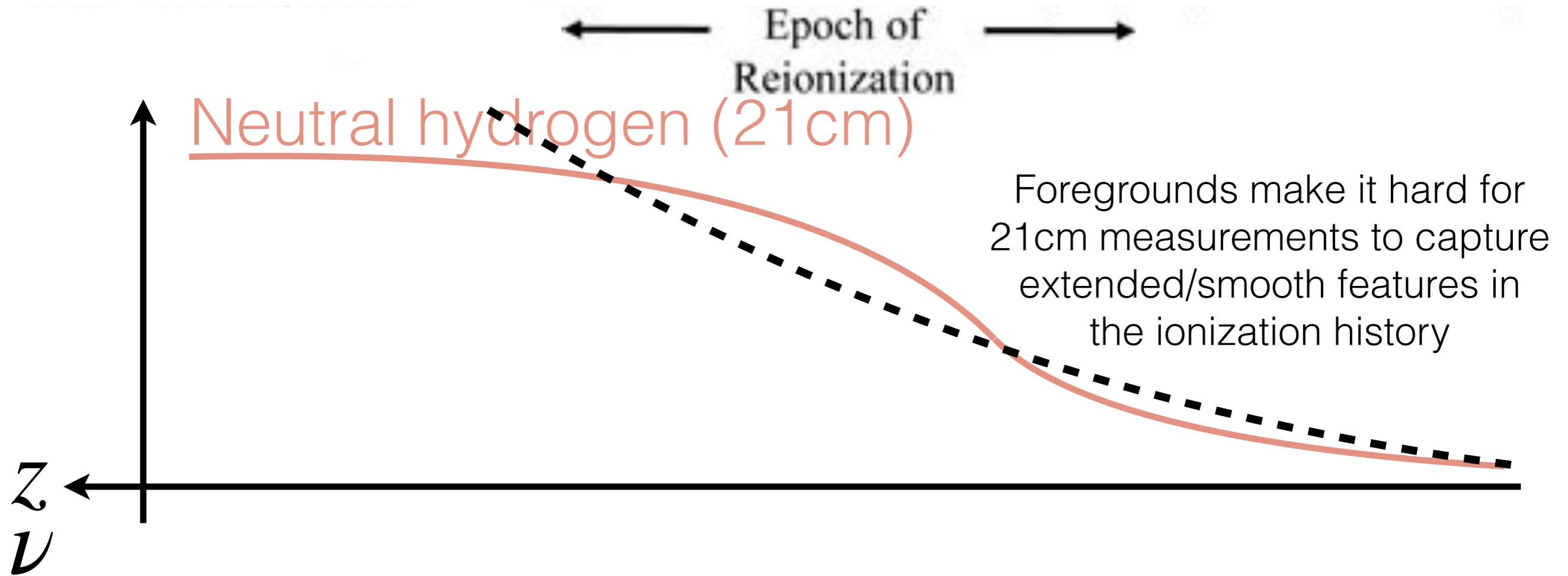
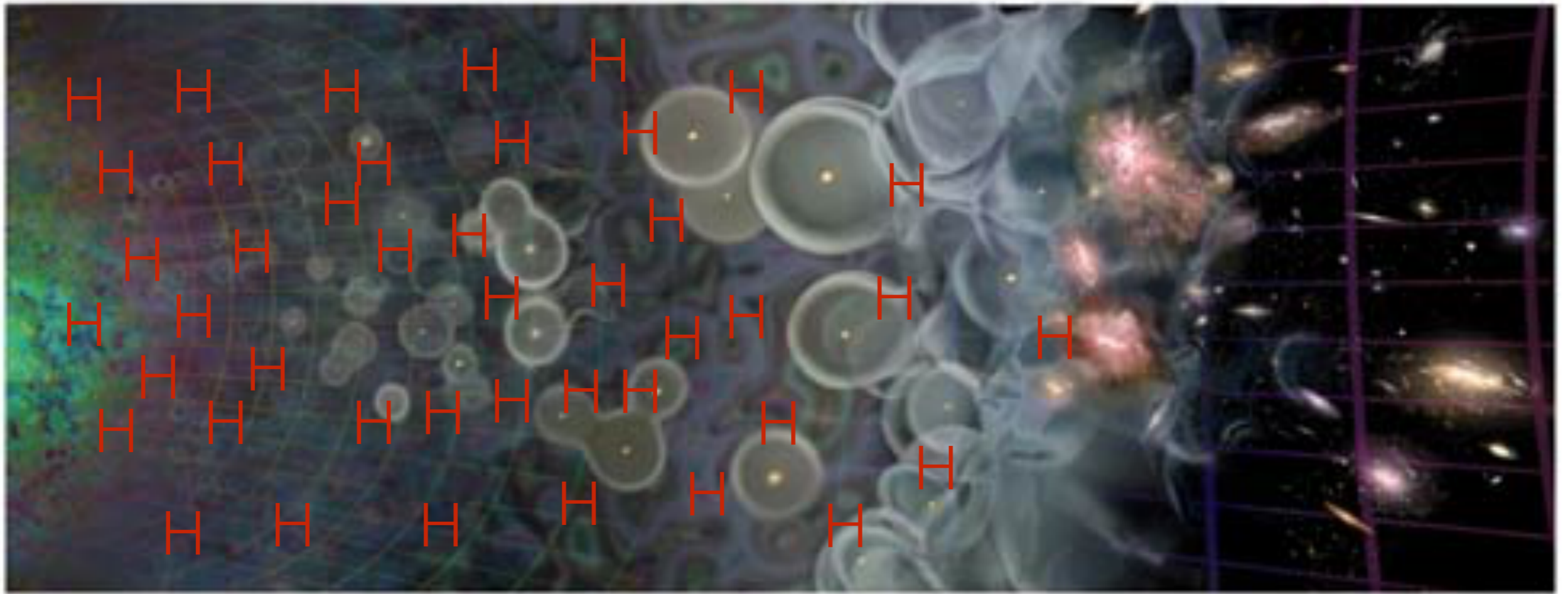


← Epoch of Reionization →

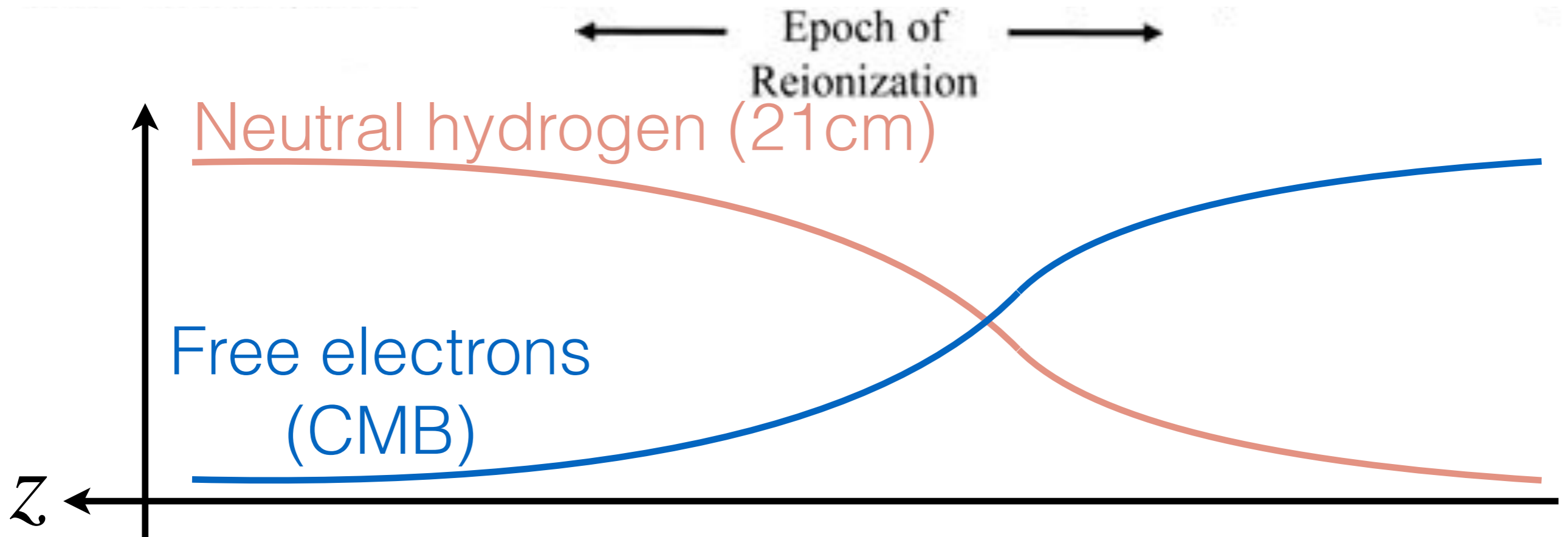
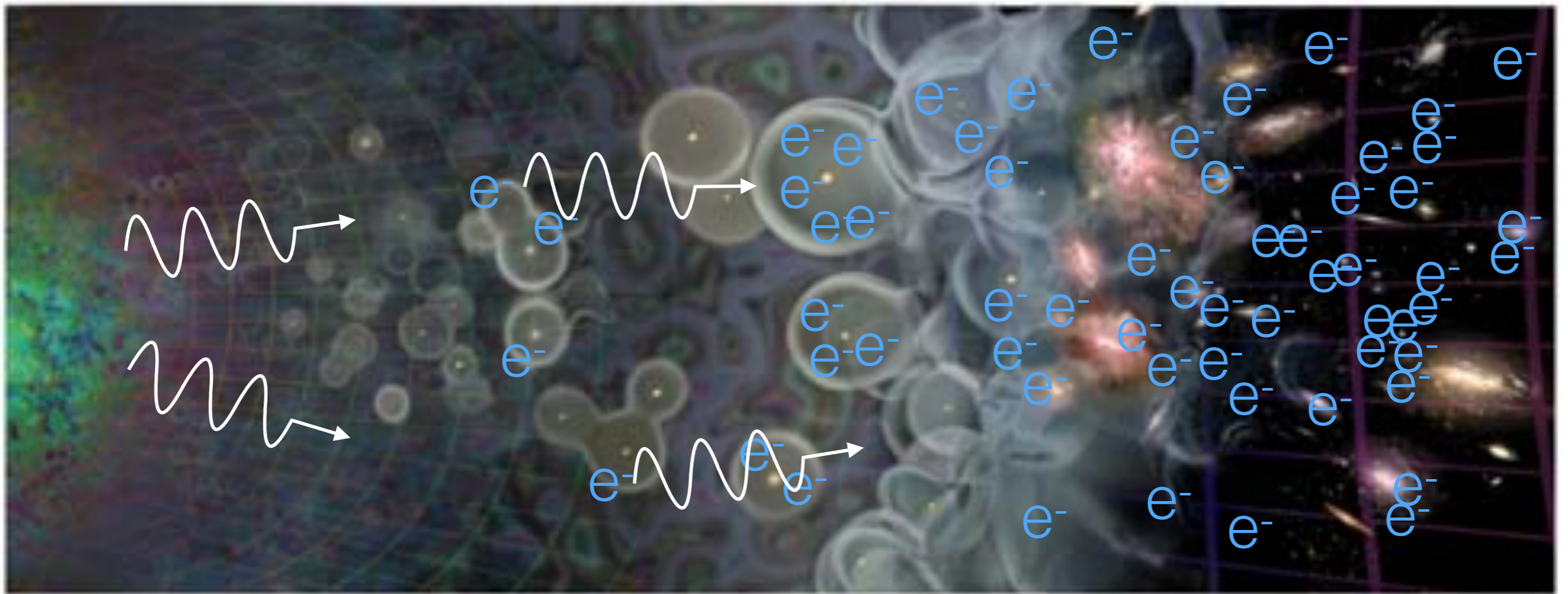
Neutral hydrogen (21cm)

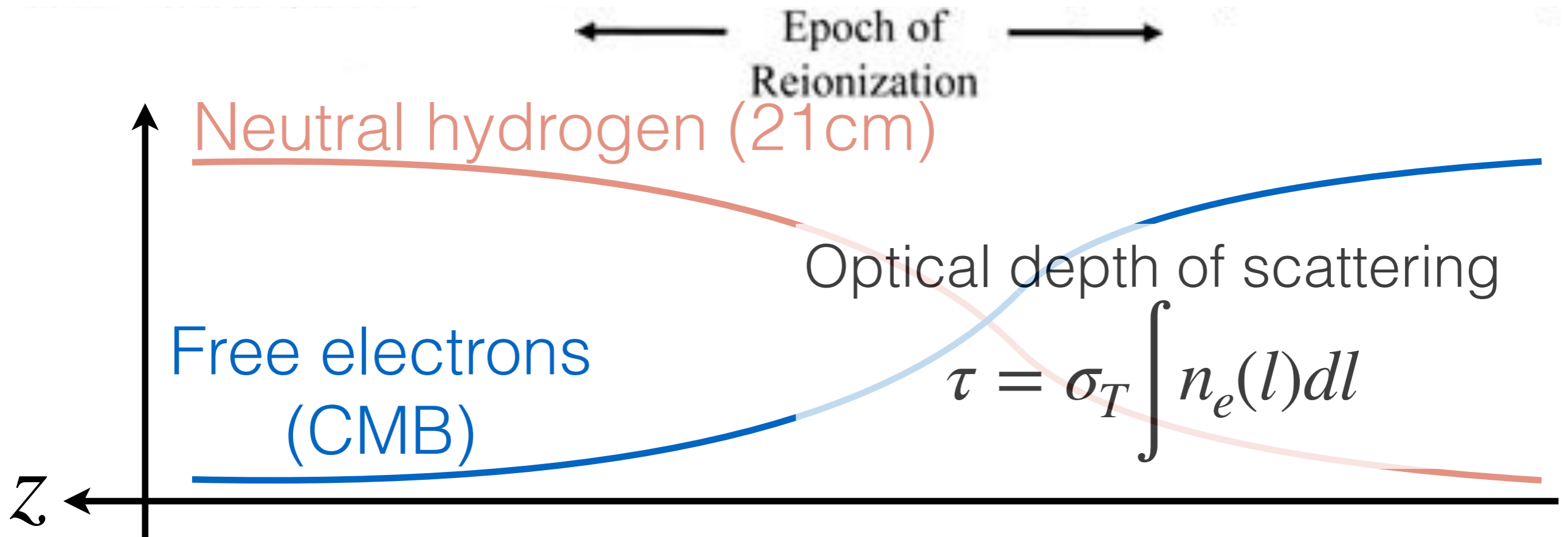
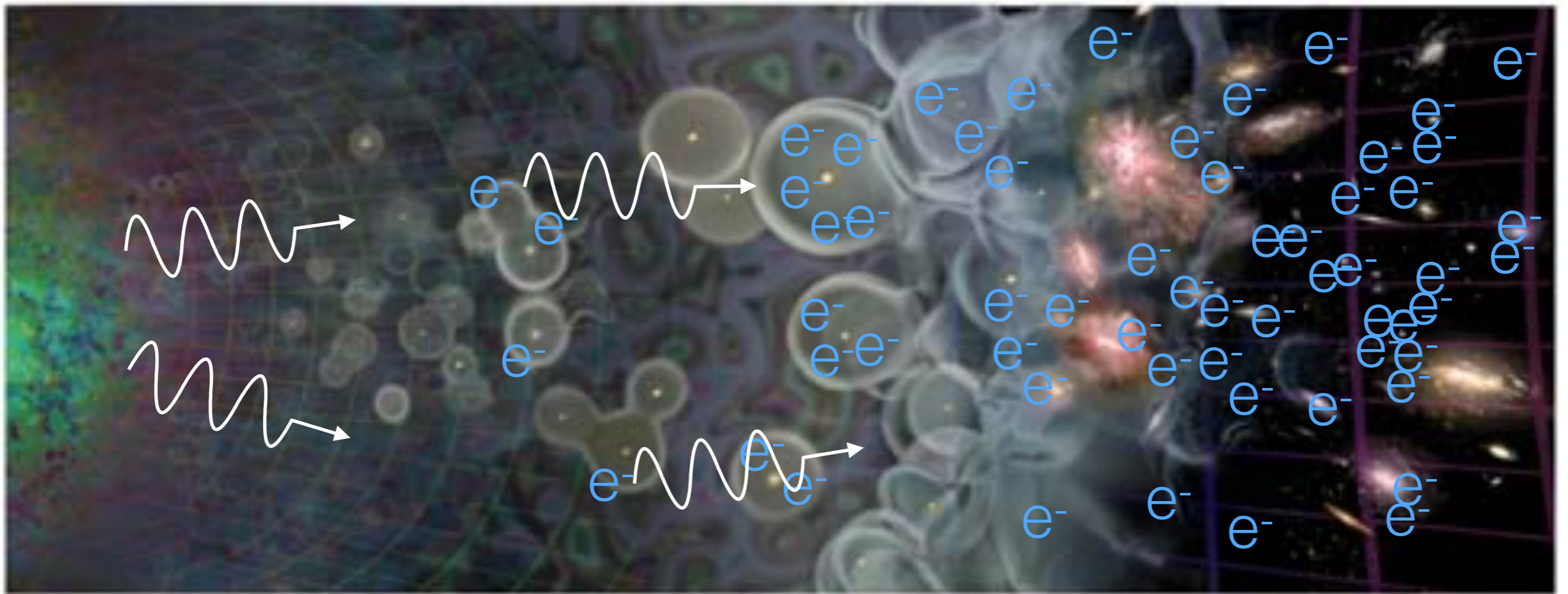
$z$

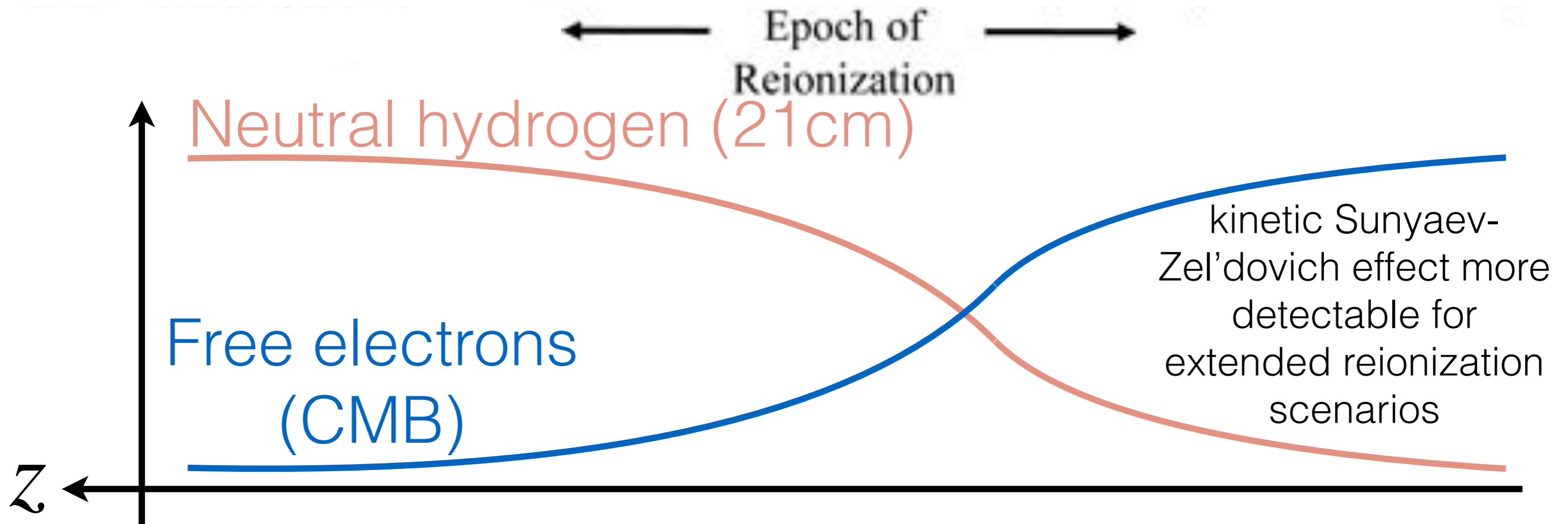
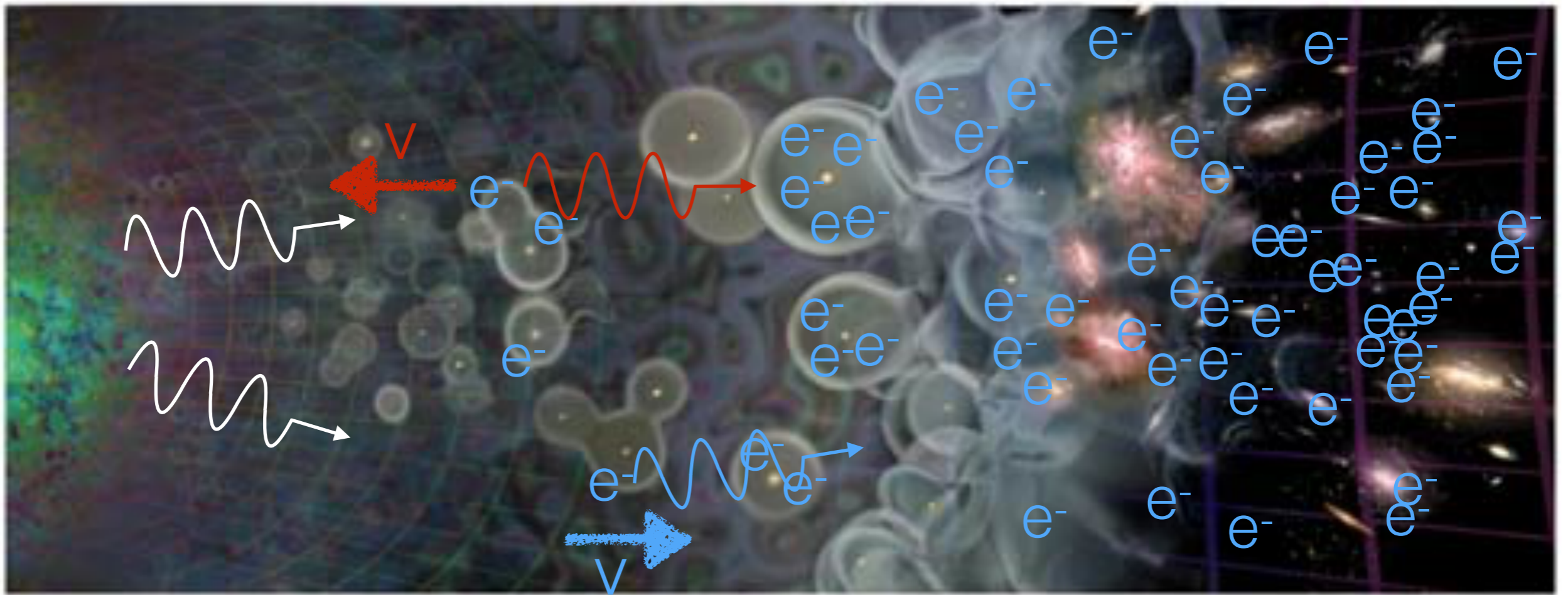












The complementarity of our two probes can be brought out using a **Karhunen-Loève basis to describe** reionization histories

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$$\mathbf{F}_{\text{kSZ}} \mathbf{v} = \lambda \mathbf{F}_{21} \mathbf{v}$$

$N_z \times N_z$  Fisher information matrix for kSZ ionization history measurements

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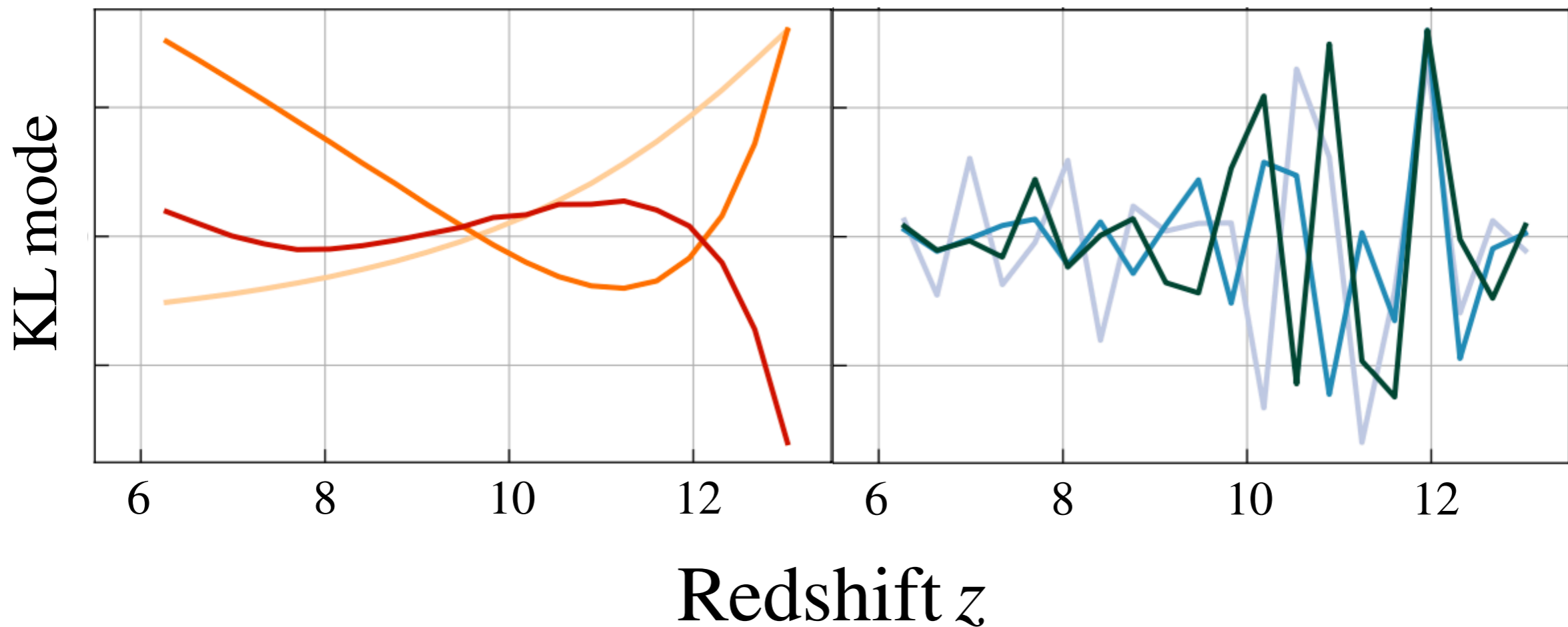
$$\mathbf{F}_{\text{kSZ}} \mathbf{v} = \lambda \mathbf{F}_{21} \mathbf{v}$$

$N_z \times N_z$  Fisher  
information matrix for  
kSZ ionization history  
measurements

Solving this generalized eigenvalue problem gives a set of basis vectors **rank-ordered by how well they can be probed by 21cm vs kSZ**

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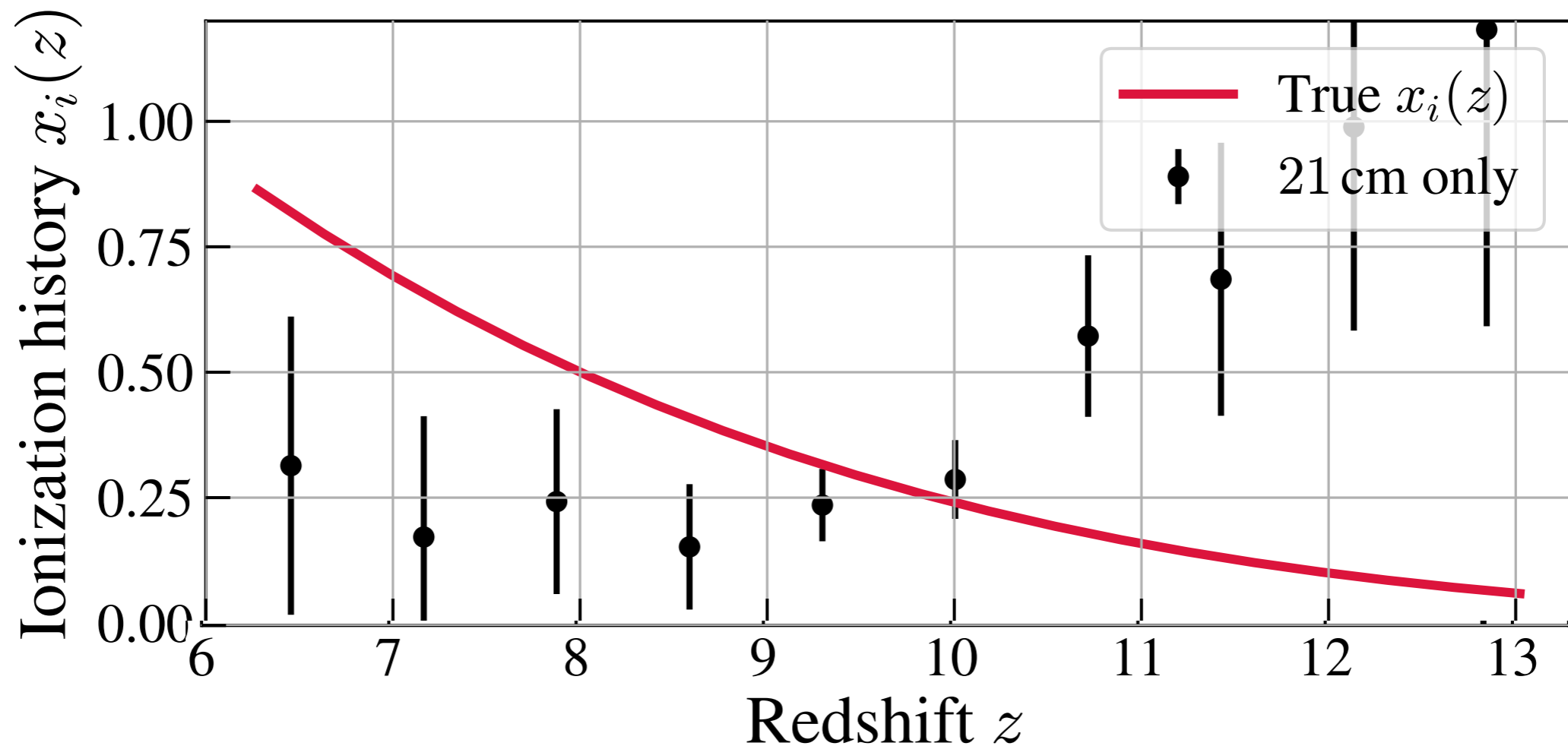
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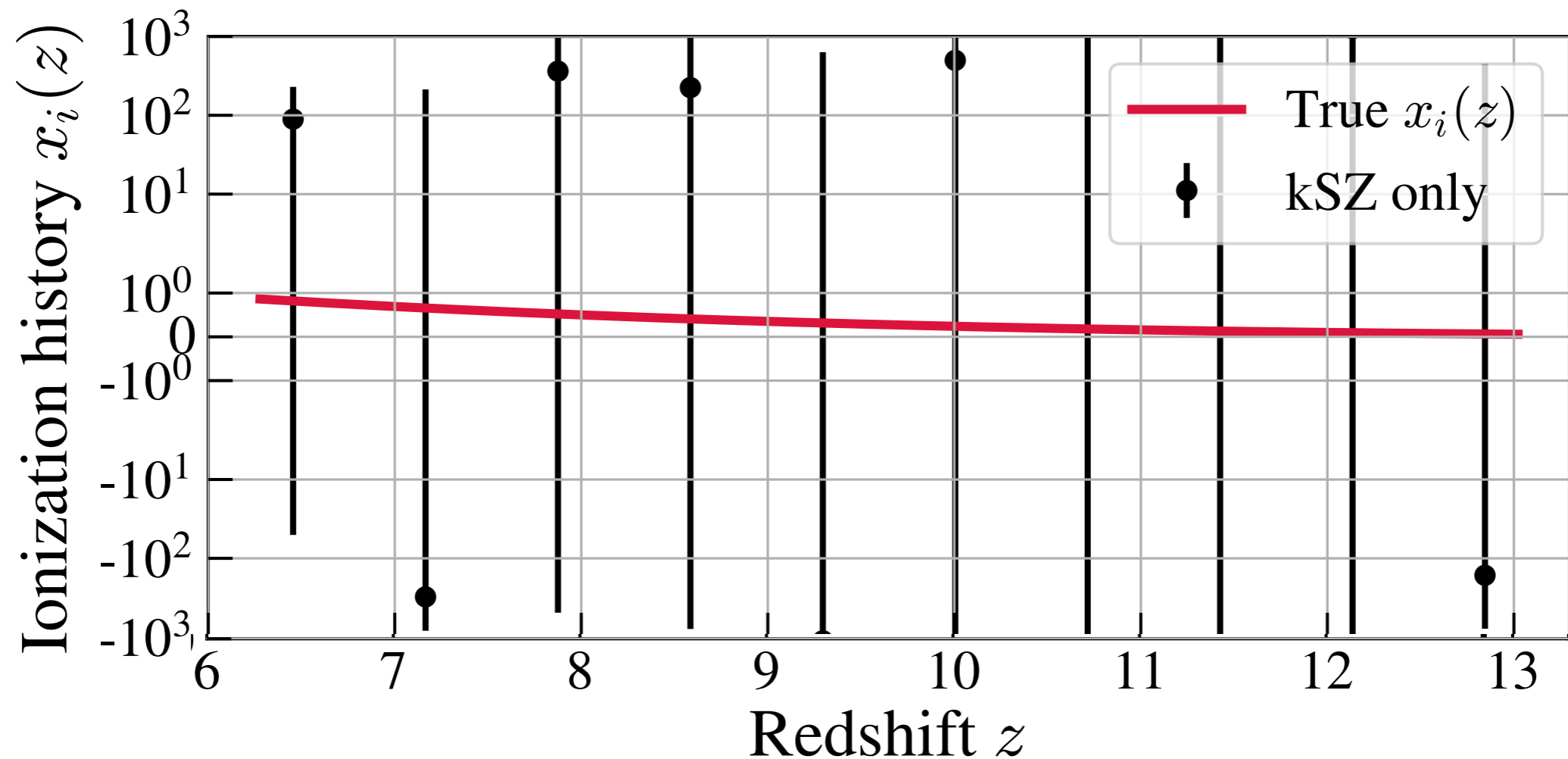
Higher-order CMB information (like from the **kinetic Sunyaev-Zel'dovich** effect) can be used in conjunction with the **global 21cm signal** to produce **model-independent constraints** on reionization histories



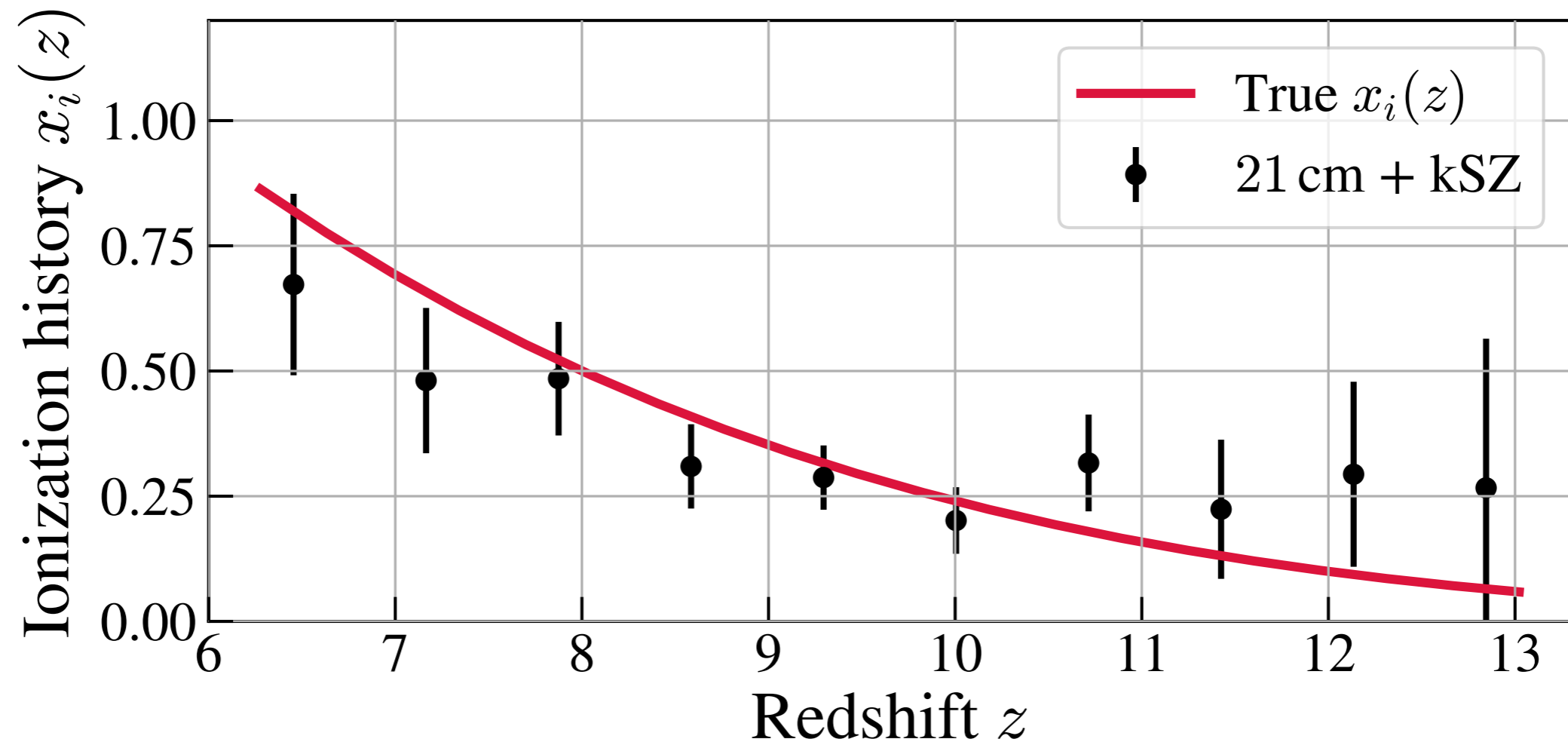
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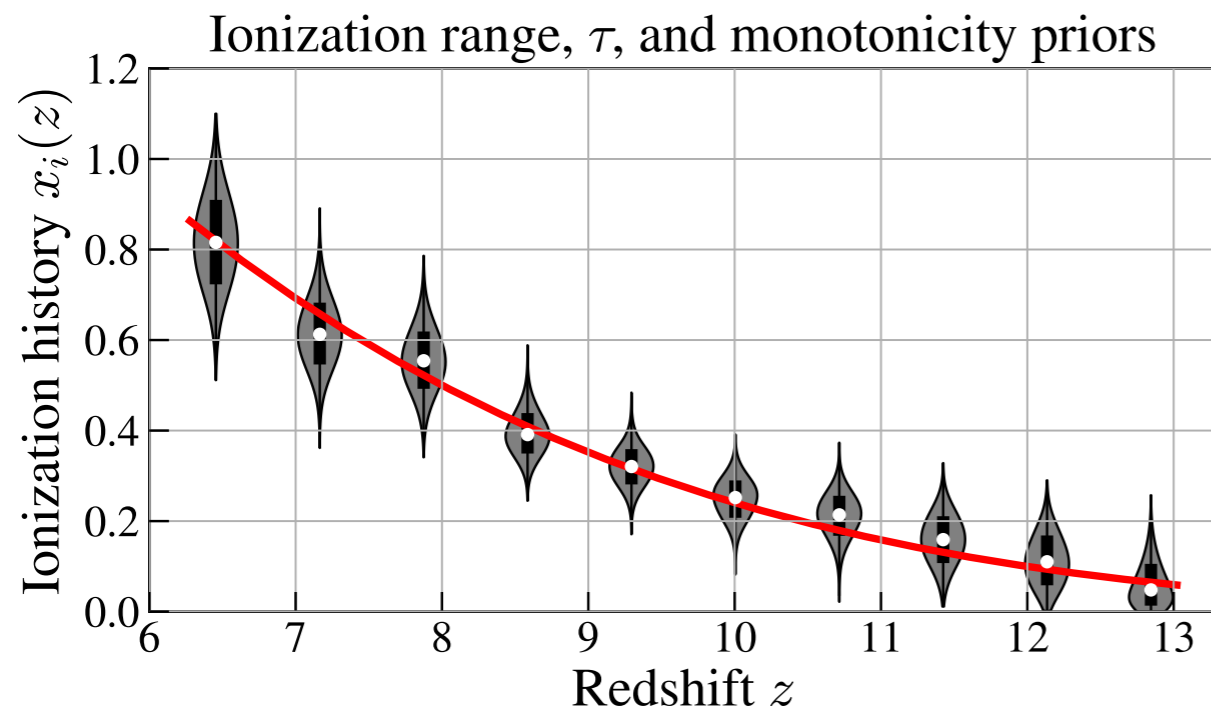
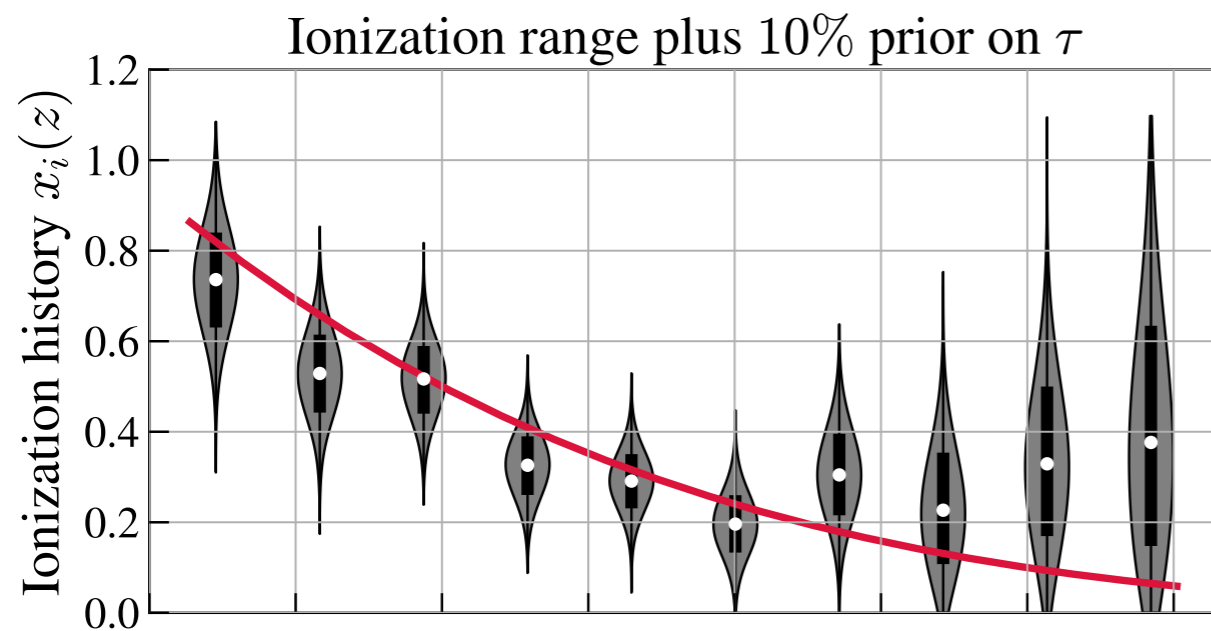
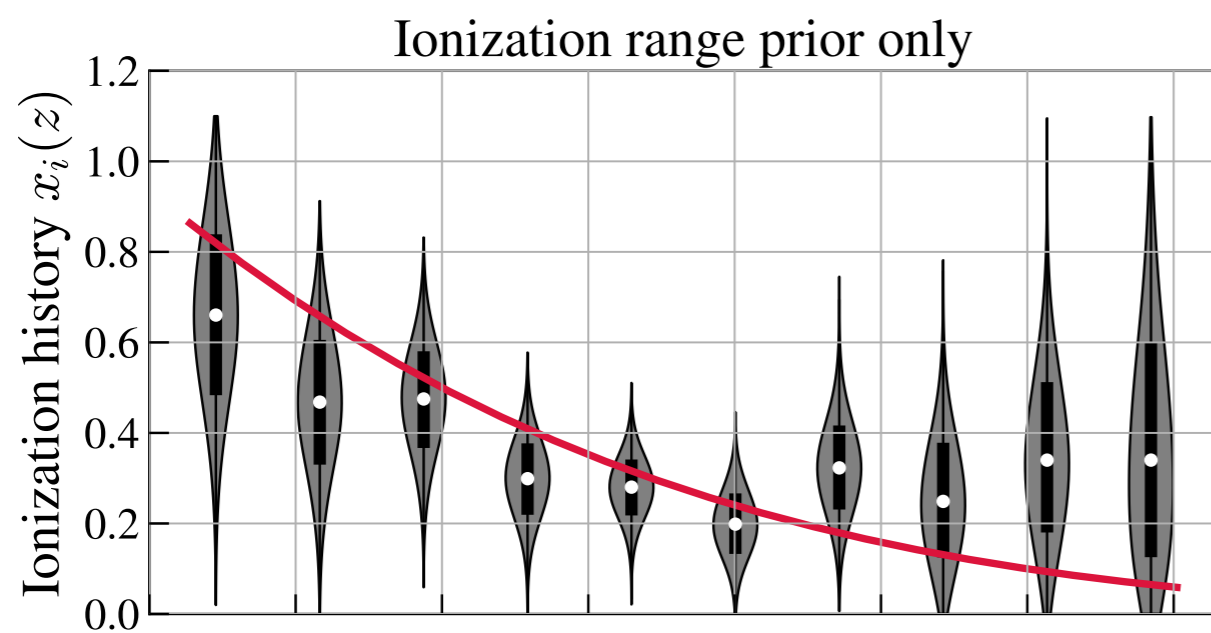
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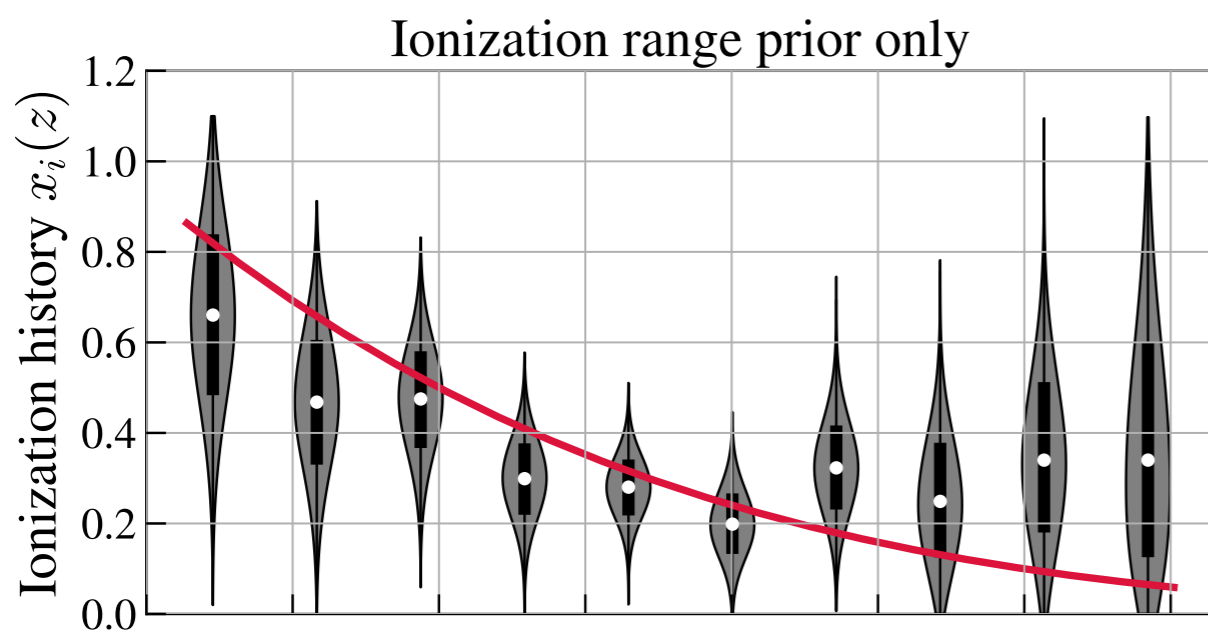


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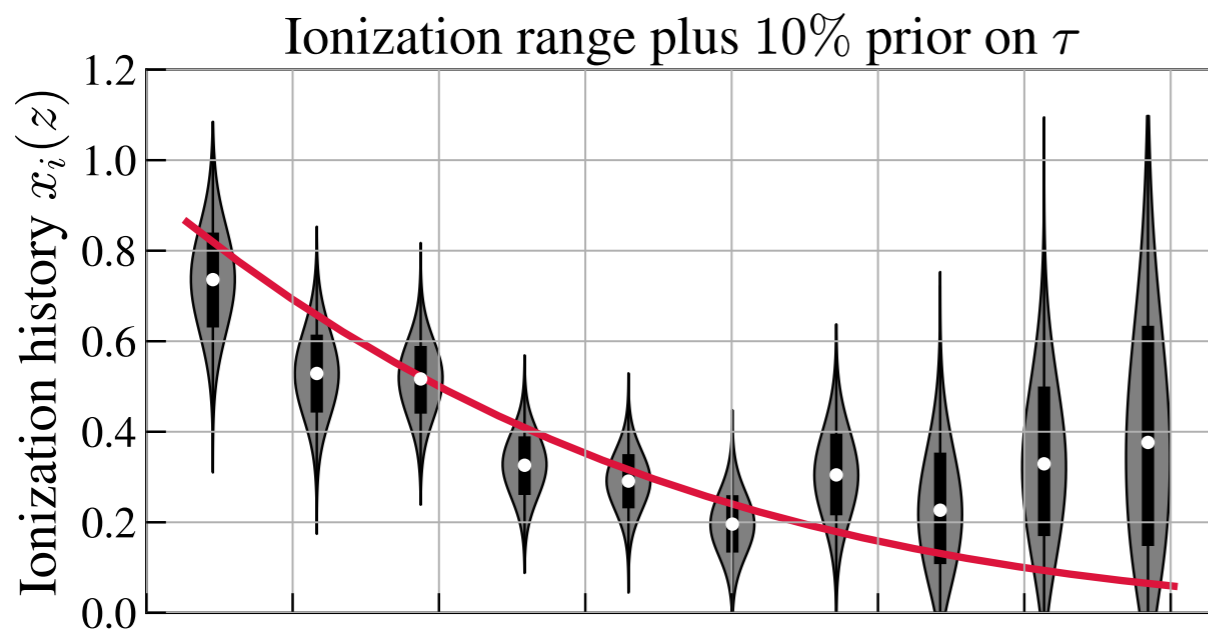


Adding other priors  
can further sharpen  
constraints....



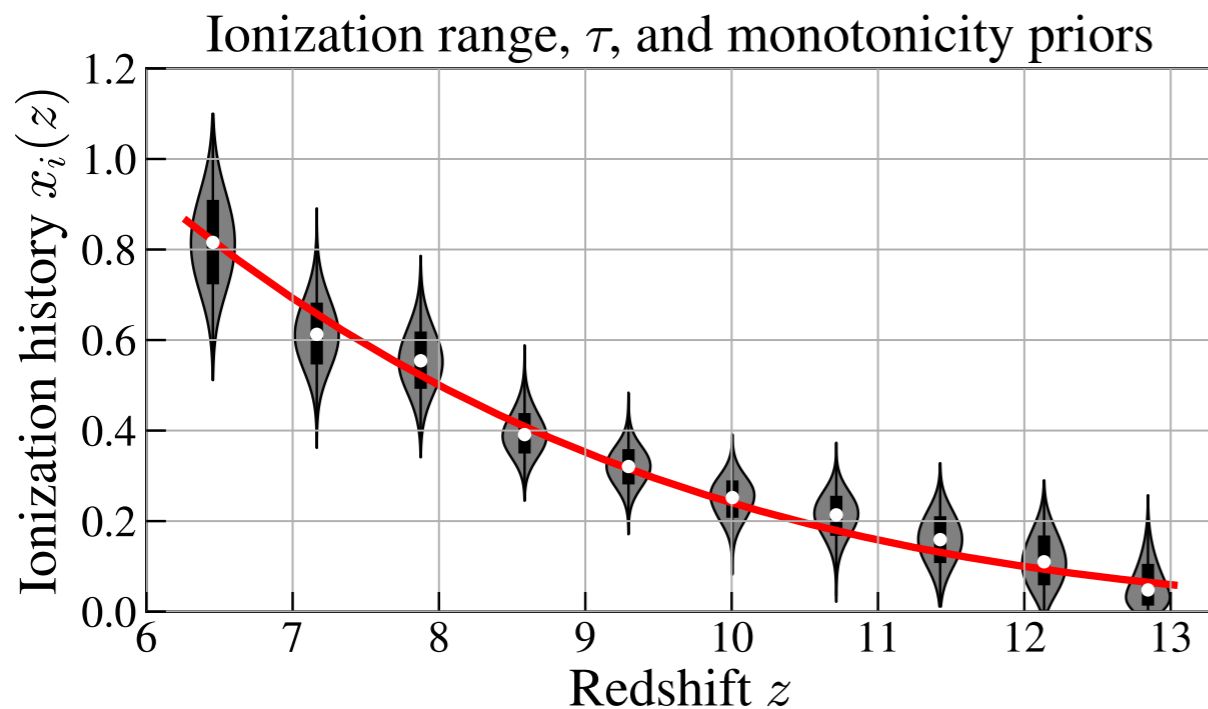


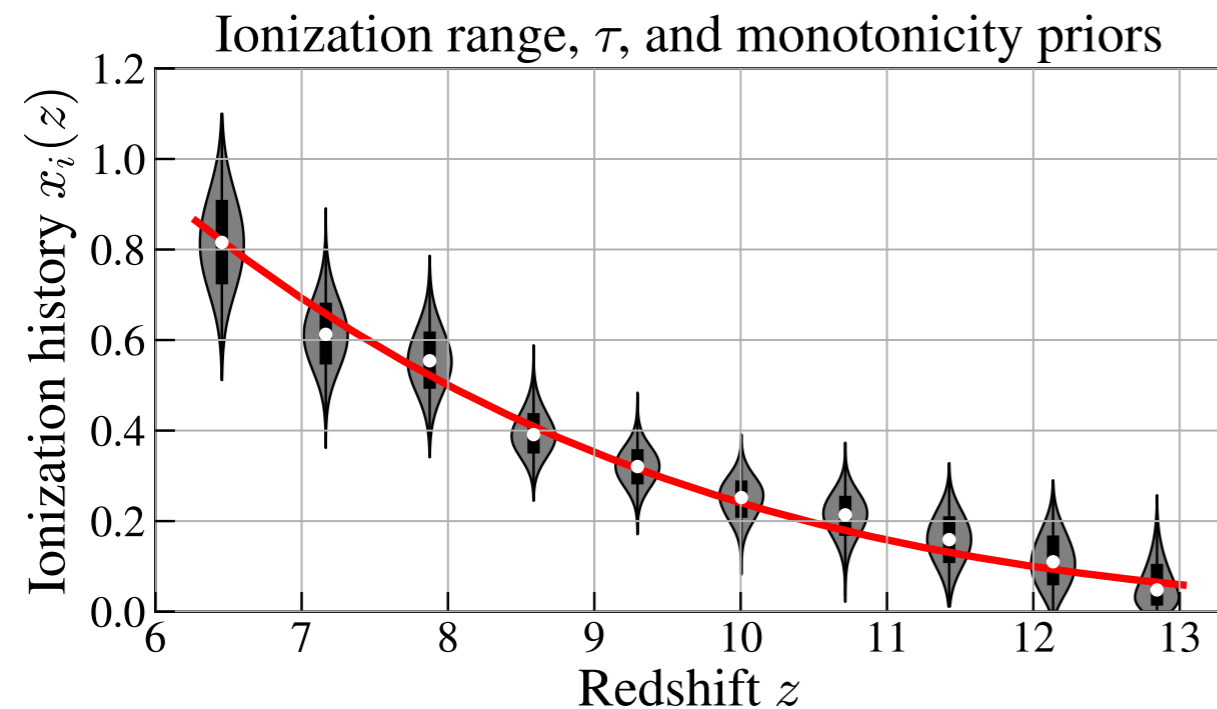
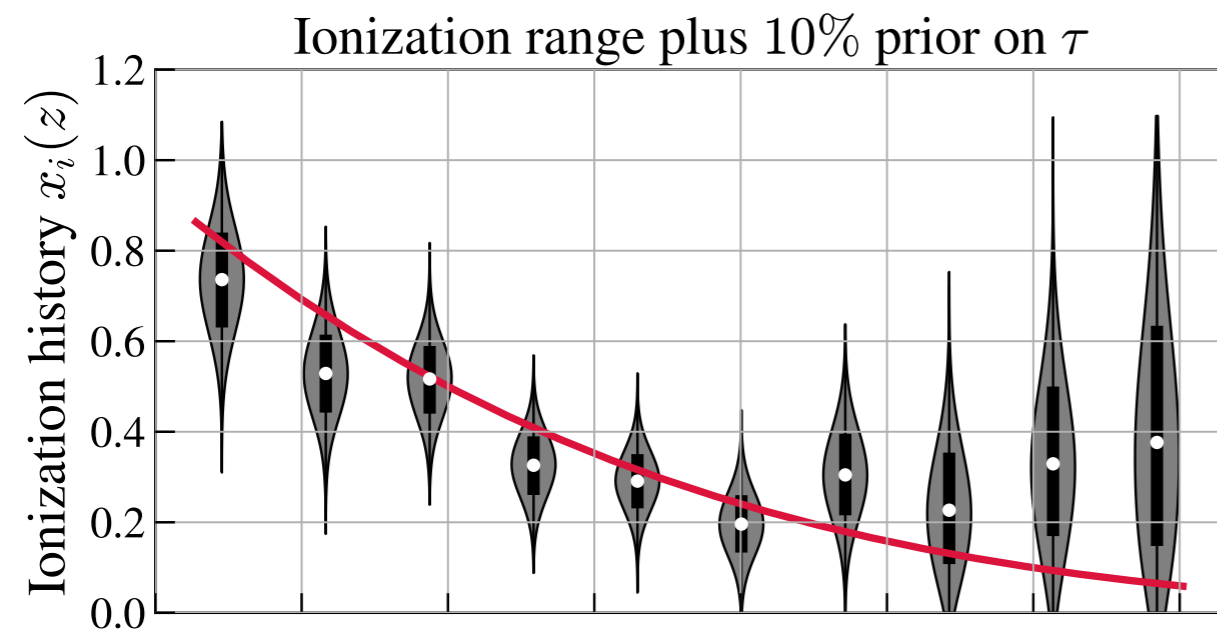
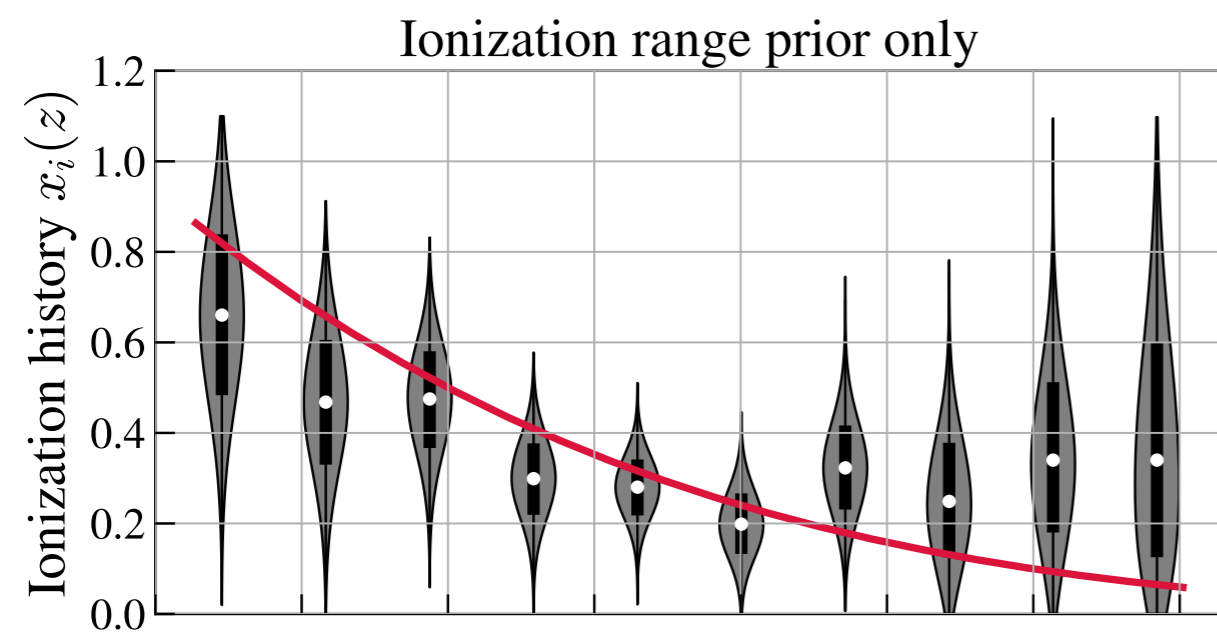
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can further sharpen  
constraints....



Optical depth of scattering

$$\tau = \sigma_T \int n_e(l) dl$$





Adding other priors  
can further sharpen  
constraints....

...and combining  
probes also  
provides safeguards  
against potential  
systematics

Bégin, **AL**, Gorce (2022),  
PRD **105**, 083503